

Ground-based direct imaging of extra-solar planets supported by AO

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Abstract. The detection of extra-solar planets made by direct imaging is an extremely challenging goal for astronomers due to the possibility to access the physical properties of planets and not only their existence. Using 8 m class telescopes joint with dedicated techniques (such as Simultaneous Differential Imaging aiming to suppress the speckle noise) it is possible, at present, to attain detection limit of 9–11 mag at 0.5 arcsec i.e. to access 3–10 M_J planets orbiting around young (100–200 Myr) nearby and late type stars. Searches for extra-solar planets carried out with the present technology are quite fundamental and critical not only for discovery of planets but also because it permits us to put constraints on theories of planets formation and migration. Besides, our understandings of the performances of sophisticated techniques such as the SDI is fundamental to plan new observational strategies, new generation instruments and telescopes. Speckles noise is, indeed, the main source of noise for observations in the NIR and visible and our ability in suppressing it is not so easily scaled at different parameters space. In this contribution I will present the main results that we obtained in on-going searches for planets carried out with NACO and NACO/SDI in the last years. A particular attention will be dedicated in comparing different observational strategies and in the employment of image processing techniques for recognizing, in an automatic way, planet features in deep images obtained with ground-based telescopes and AO facilities.

Keywords. planetary systems, image processing.

1. Introduction

Adaptive optics (AO) imaging techniques at 8–10 meter telescopes can detect young planets if they are sufficiently distant from their parent star. Since it is reasonable to assume a coevality between the parent star and the planet, we can retrieve from atmospheric models (Burrows *et al.* 1997, Baraffe *et al.* 2003) that planets having a mass in the range 3–10 M_J and orbiting around young late-type stars (10–200 Myr) have a typical brightness contrast with respect to the parent star of the order of 10^2 – 10^6 , corresponding to a magnitude difference of ΔM of 5–15 mag. This means that a planet can be detected at a few tens of AUs from parent stars having distances of ~ 50 pc.

According to recent statistical results (e.g. Marcy *et al.* 2003), most of known exoplanets orbit at distances smaller than 1 AU and the semi-major axes of the exoplanets found so far are not larger than ~ 6 AU. Since these results are heavily biased towards short separations the derived major-axis distribution is certainly not representative. Not much is known at present about potential planets that could exist at distances larger than about 6 AU. Besides, there are models (for example the gravitational scattering model) that claim that in multi-planetary systems, due to gravitational interactions between the

planets, it can happen that the lightest objects are ejected on hyperbolic trajectories and some planets can move on stable orbits having a distance of a few tens AU from the parent star. Such models were proposed by Weidenschilling & Marzari (1996) and Rasio & Ford (1996) and studied more recently by Papaloizou & Terquem (2001).

There are thus elements that indicate that such planets might exist at large distances from the central star. For this reason we planned and performed an extended survey aiming to search for massive exo-planets orbiting around nearby and young stars using one of the best available instruments at the present time for high-contrast AO imaging observations: NACO at the Very Large Telescope (VLT). Once the survey completed (Masciadri *et al.* 2005), we decided a follow-up of our project employing the SDI/NACO instrument. As explained in Section 3, such an instrument permits to improve the contrast reachable at sub arcsecond distances from the parent star. I report here the principal results obtained and I discuss the performances of the different observational strategies. Finally, I deal about a method that a few of us (Masciadri & Raga, 2004) developed for recognizing in an automatic way, planets features in deep images.

2. Deep direct imaging survey

We refer to Masciadri *et al.* (2005) for detailed informations on the target selection, observational strategy employed in the survey and data processing. Here we briefly summarize the principal informations. Around 30 targets (nearby stars) were observed during 3 runs. The targets are placed at a distance $d \leq 77$ pc (with 70 % of them in the first 32 pc), they have an age of 1–200 Myr and late spectral type (K and M). Firstly we observed each target in a narrow band filter for a few minutes in order to obtain a non-saturated point spread function (PSF). Secondly, we obtained deep (~ 20 minutes) K_s and H broadband observations of each target. Data were reduced using standard procedures to eliminate bad pixels, flat and background sky (IRAF and ECLIPSE). A dedicated pipeline has been employed to filter the low spatial frequencies in order to improve the detection limit at subarcsec distances (Masciadri *et al.* (2005)). Two different methods were used for the low spatial frequencies filtering.

We averaged the contrast ΔM obtained at 0.5 and 1 arcsec for all the observed targets having a total integration time within the range 20–30 minutes. We find, in the K_s band, a $\Delta M = 9$ mag at 0.5 arcsec and a $\Delta M = 11.5$ mag at 1 arcsec. We find, in H band, a $\Delta M = 9.2$ mag at 0.5 arcsec and a $\Delta M = 11.7$ mag at 1 arcsec. We recall that 0.5 arcsec corresponds, for our sample, to projected separations of 2–35 AU, 1 arcsec to projected separations of 4–70 AU. This illustrates nicely that NACO is a well performing instrument for high-contrast imaging observations.

No very promising massive-planet candidate was identified in our survey. Besides this we calculated the detection limit for 28 of these targets (Masciadri *et al.* 2005) and we retrieve the upper limit of the star/planet projected separation calculated for a potential 10 and 5 M_J planet (Masciadri *et al.* 2005, Table 4). By analysing these results from a statistical point of view, we can calculate (Masciadri *et al.* 2005, Fig.9) the cumulative distribution of the upper limit of the distance star/exo-planet for our sample of targets. The thin line represents the cumulative distribution obtained for 5 M_J planets. We find that, with respect to our sample of targets, in 50% of cases (median value), there are not planets at distances larger than 14 AU and, in 100% of cases, there are no planets at distances larger than 65 AU. We underline that we are considering a 5 M_J planet just as a reference, but the same calculation could be done for more and/or less massive planets

in the range 3–10 M_J . In the case of more massive planets, the cumulative distribution should have a medium value (50%) associated to an angular separation smaller than 14 AU. In the case of less massive planets, the cumulative distribution should have a medium value (50 %) associated to a angular separation larger than 14 AU. To better appreciate the sensitivity with respect to the mass of the planets we also calculated the cumulative distribution for 10 M_J planets (Fig.4, bold style line). We find that, in 50% of the cases (median value) there are no planets at distances larger than 8.5 AU and in 100% of the cases there are not planets at distances larger than 36 AU. The number given in Table 4 (Masciadri *et al.* 2005) and the cumulative distribution illustrate that it is relatively unlikely that a massive planet exist beyond a certain distance from the star. For example the fact that no 5 M_J planet was found for 18 of our targets beyond 20 AU (see Table 4, Masciadri *et al.* 2005) means that the likelihood for the existence of such planets beyond this distance is $\sim(1/18)$ % i.e. $\sim 5.6\%$. This last estimation is useful for the reader to have an idea of how much the size of the sample affects the statistical results. We remind that we are using the term *probability* in a less conventional way. A precise (statistically speaking) definition of *probability* should need some positive detection that we obviously do not have.

3. Contrast at sub-arcsec distances: comparison of different techniques

In the survey carried out with NACO (Masciadri *et al.* 2005) no interesting candidates were detected but typical contrasts at 0.5 arcsec and 1 arcsec were calculated using the detection limit of 28 targets in H and K_s bands (see previous Section). The success of the search for extrasolar planets strictly depends on how efficient we are in increasing the reachable contrast at sub-arcsecond distances. We calculated, therefore, for deep direct images, the contrast vs. angular separation for different exposure times. Figure 1 (left) shows the result of this calculation. At distances larger than 1 arcsec, the larger is the exposure time, the larger is the reachable contrast. At sub-arcsec distances, the increasing of the exposure time does not supply a better contrast because we are fundamentally limited by the speckle noise. The Simultaneous Differential Technique (SDI) proposed by Racine *et al.* (1999) is a rather efficient method to reduce the speckle noise at small angular separations from the central star. Images taken simultaneously at slightly different wavelengths are subtracted from each other. Since the speckle noise of both images is nearly the same[†], the speckle noise is eliminated in the resulting image providing one can sufficiently limit static aberrations. The differential technique finds a useful application in detecting cool methane-rich faint objects. Rosenthal *et al.* (1996) first proposed observations of the same star done simultaneously in two different wavelengths which are located “on” and “off” the methane (CH_4) absorption band. If one subtracts the image taken in the absorption band from the one taken outside the absorption band, planet/brown dwarf features appear in the continuum. The absorption at 1.62 μm , in the CH_4 band, is a clear signature of any object having temperature $T \leq 1300^\circ \text{K}$ (giant exo-planets and T-type brown dwarfs). Supported by the preliminary results obtained by Close *et al.* (2004) with the commissioning data of the SDI/NACO instrument we decided to start a new survey in collaboration with the Close group (see Biller contribution, this Proceedings). In this contribution we intend to study the efficiency of the different observational strategies at sub-arcseconds distances from the target to estimate the improvement obtained with the

[†] We precise that speckles of images taken at slightly different wavelengths match if frames are rescaled opportunely by a λ_2/λ_1 factor.

Table 1. Mean contrast at 0.5 arcsec calculated with different observational techniques (deep direct imaging and SDI). ^(a)We remind to the reader that the SDI technique is employed in narrow band.

D (m)	Obs. Tech.	Filter	Δm (mag)	$N^\circ\sigma$	Exp. Time (minutes)	N° Targets
8	Deep Imaging	NACO	H	9.20	5σ	20-30
8	SDI/NACO ^(a)	$1.57\mu\text{m}$	H	8.61	5σ	10-20
8	SDI/NACO	H	10.31	5σ	10-20	15

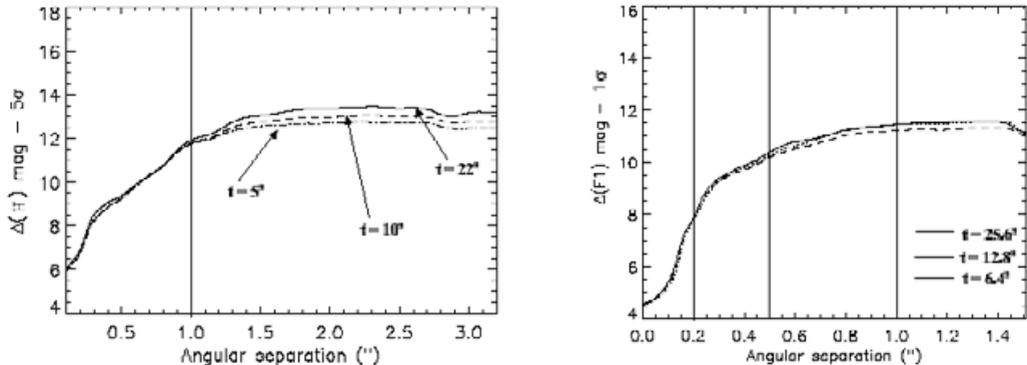


Figure 1. Left: detection limit vs. angular separation of one target of the survey done with deep direct imaging (Masciadri *et al.* 2005) obtained with integration times of 5, 10 and 22 minutes. Right: detection limit obtained with the SDI technique and a total exposure time of 6.4, 12.8 and 25.6 minutes.

SDI on sky on a 8 m class telescope. Table 1 summarize the mean contrast calculated at 0.5 arcsec employing the deep direct imaging and the SDI technique running with the same instrument (NACO). We want to precise that, employing the SDI technique, the calculation of the contrast refers to the narrow band, more precisely the narrow band at $1.57 \mu\text{m}$ ($\text{FWHM} = 0.024\mu\text{m}$). In order to compare the two observational strategies (Direct deep imaging and SDI) it is necessary to calculate the equivalent contrast in H broadband. This means that $\Delta m_H = \Delta m_{1.57\mu\text{m}} + A$. The factor A depends on the spectral type of the planet-like object because the spectrum shape in proximity of the absorption band is different for different spectral type planet-candidates. In the case of a T6 type object we calculated that $A = 1.7 \text{ mag}$. For spectral type later than T6, A is larger than 1.7 mag, for spectral type earlier than T6, A is smaller than 1.7 mag.

Looking at Table 1 we conclude that the **SDI technique permits to gain a factor of 1 mag in H band with respect to the Direct deep imaging at 0.5 arcsec**. We underline that this gain is a typical value, obtained averaging detection limit curves related to 15 targets and it is valid for the SDI applied to a 8 m telescope. Figure 1 (right) shows the detection limit at $F_1 = 1.57 \mu\text{m}$ obtained for different exposure time: 6.4, 12.8 and 25.6 minutes. One can observe that, in the range 0.5–1 arcsec the gain is $\Delta F_1 \sim 0.5 \text{ mag}$ in $\Delta t = 15 \text{ minutes}$. In the range 0.2–0.5 arcsec, the gain ΔF_1 is $\sim 0 \text{ mag}$ in the same Δt . This means that, at least in the actual version, the gain of the SDI is weakly depending on the total exposure time.

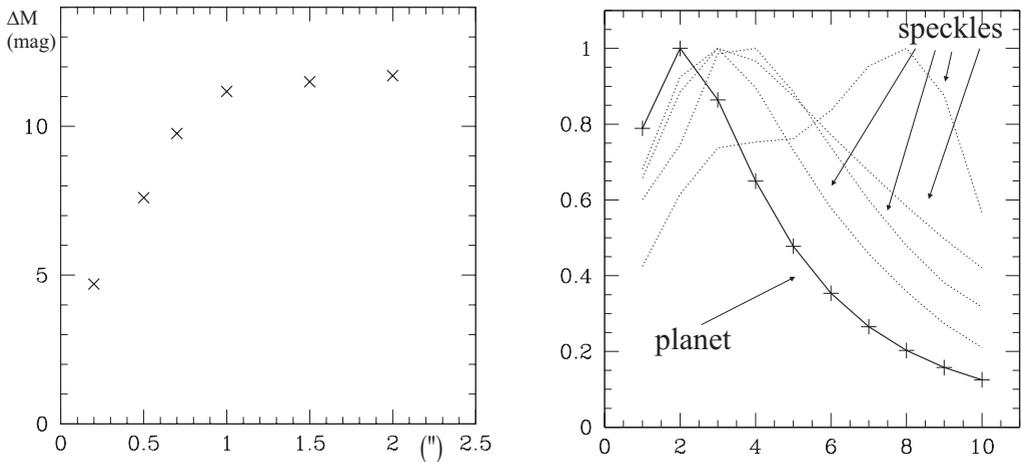


Figure 2. Left: ΔH (mag) between the central star and the template planet recognized by the automatic procedure vs. the angular separation (arcsec). Right: Example of a normalized spectrum of a planet (y-axis) vs. number of pixels (x-axis). This spectrum is used as template spectrum to recognize planets in a deep image. Dotted lines: spectra of a few speckles recognized by the procedure as local maxima and rejected because they did not follow the criteria of selection.

4. Automatic recognition of sub-stellar PSFs with wavelet technique

The fundamental problem for planets detection done with the direct imaging technique is that we have a huge contrast at very small angular separations. It would be therefore quite useful to develop a procedure, better if this is automatic, for recognition of features typical of sub-stellar PSFs. The method that we proposed (Masciadri & Raga 2005) is based on the wavelet technique. The detailed description of such a procedure can be found in Masciadri & Raga (2005). The principle is summarized here: a template planet, obtained scaling opportunely the no-saturated PSF of the star, is added in the deep image of the central star. The resulting deep image is filtered from the low spatial frequencies. The procedure characterizes the spectrum of the template planet (Fig.2-right, solid line) and then it looks for, in the deep image, features having the same spectra. These features have to be considered as potential sub-stellar or stellar-like PSF. A second epoch of observation is, therefore, necessary to discriminate between the two cases. This procedure was recently validated (Masciadri & Raga 2005) on a target (V343 Nor, $m_k=6.3$ mag, $d \sim 40$ pc, age ~ 12 Myr) of our NACO survey. Figure 2-right shows the contrast (ΔH) between the central star and the template planet versus the angular separation within the 0.2-2 arcsec range at which a planet is recognized by the automatic procedure. We note that, in the case of V343 Nor, a maximum $\Delta H = 4.7$ mag, associated to a planet, can be recognized in an automatic way at 0.2 arcsec with the present procedure. To give an idea to the reader of what we are talking about, we note that 4.7 mag is roughly the contrast in K band of 2M1207, the giant planet candidate recently discovered by Chauvin *et al.* (2004), with respect to its central star. Figure 2-right shows the spectra of a few “speckles” (dotted lines) recognized by the procedure as local intensity maxima that are rejected by the procedure because their spectra do not follow the criteria of selection (see ‘tolerance criterium’ in Masciadri & Raga 2005). The spectrum shown in bold style (Fig.2-left) refers to the template planet placed in the deep image. One can observe that the procedure can discriminate the spectra of planets and speckles in an efficient way.

5. Conclusions

We can summarize the main results obtained in this study in the following way:

- $5M_J$ planets, having an age within 200 Myr, orbiting around stars placed at a distances $d \leq 70$ pc can be detected with NACO at distances smaller than 10 AU.
- We find a low probability ($\sim 10\%$) to find $5M_J$ planets at distances larger than 26 AU and $10M_J$ at distances larger than 16 AU around stars with properties defined above.
- The performances of different techniques (running at present on sky) at sub-arcsecond distances can be summarized in the following way. At 0.5 arcsec, the SDI technique permits to reach a typical contrast at 5σ , in [10-25] minutes, equal to $\Delta H = 10.31$ mag. This is around 1 magnitude better than what obtained with the deep direct imaging ($\Delta H = 9.2$ mag). Results published recently (Boccaletti *et al.* 2004) indicate that the 4 quadrant phase mask (4QPM) coronagraph implemented on NACO gives, at 0.5 arcseconds, a $\Delta H = 9$ mag at 5σ in K band. In order to do a fair comparison with the deep imaging and the SDI technique, results obtained in H band should be necessary. Considering that the SDI and the 4QPM are certainly techniques that will support the high-contrast future facilities of 8 m class telescopes and ELTs, we think that it is quite fundamental to verify on sky the performances of these techniques to be able to know where and how to improve such techniques and to scale performances at different parameters space. At present we can state that the SDI can be considered a success giving access to an improvement of the contrast at sub-arcsecond distances of around 1 mag at 0.5 arcsecond with respect to the deep direct imaging observational strategy. On the other side, we have to note that, no planet candidates have been observed so far with the SDI technique that have not been also observed with the simpler deep direct imaging strategy. From an astrophysical point of view, it would be for sure interesting to be able to do that to optimize the scientific impact of such a promising technique.
- The wavelet technique is certainly a promising method for recognizing, in an automatic way, sub-stellar planet features in deep images.

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