electron density of about 350 ± 150 cm^{-3}. Seeing limited (O III) imagery of this planetary nebula is also presented. (This article is appearing in Astron. Astrophys. in the near future).

BIPOLAR NEBULAE AND TYPE I PLANETARY NEBULAE

N. Calvet
Centro de Investigación de Astronomía, Mérida, Venezuela

M. Peimbert
Instituto de Astronomía, Universidad Nacional Autónoma de México

It is suggested that the bipolar nature of PN of Type I can be explained in terms of their relatively massive progenitors ($M_i \geq 2.4\ M_\odot$), that had to lose an appreciable fraction of their mass and angular momentum during their planetary nebula stage. The following objects are discussed in relation with this suggestion: NGC 6302, NGC 2346, NGC 2440, CRL 618, Mz-3 and M2-9. It is found that CRL 618 is overabundant in N/O by a factor of 5-10 compared with the Orion Nebula.

WIND-BLANKETED STELLAR ATMOSPHERES

D.C. Abbott, D.G. Hummer
Joint Institute for Laboratory Astrophysics

Radiation scattered by a stellar wind back into the photosphere alters the temperature-depth relation and thus the stellar flux distribution. The fraction of the radiation returned to the star at every wavelength has been calculated using stellar wind models accounting for approximately 10 000 lines. Model stellar atmospheres containing hydrogen and helium, both with and without the assumption of LTE, have been computed allowing for the reflected radiation. For realistic wind and stellar parameters relevant to central stars of planetary nebulae, we obtain a 25% increase in the surface temperature and in the optical brightness temperature, and a 2 order-of-magnitude increase in the flux in the He II continuum.