

Highly-luminous Cool Core Clusters of Galaxies: Mechanically-driven or Radiatively-driven AGN?

Julie Hlavacek-Larrondo^{1*} and Andy Fabian¹

¹Institute of Astronomy, University of Cambridge,
Madingley Road, Cambridge CB3 0HA
email: juliehl@ast.cam.ac.uk

Abstract. Cool core clusters of galaxies require strong feedback from their central AGN to offset cooling. We present a study of strong cool core, highly-luminous (most with $L_X \geq 10^{45}$ erg s⁻¹), clusters of galaxies in which the mean central AGN jet power must be very high yet no central point X-ray source is detected. Using the unique spatial resolution of *Chandra*, a sample of 13 clusters is analysed, including A1835, A2204, and one of the most massive cool core clusters, RXCJ1504.1-0248. All of the central galaxies host a radio source, indicating an active nucleus, and no obvious X-ray point source. For all clusters in the sample, the nucleus has an X-ray bolometric luminosity below 2 per cent of that of the entire cluster. We investigate how these clusters can have such strong X-ray luminosities, short radiative cooling-times of the inner intracluster gas requiring strong energy feedback to counterbalance that cooling, and yet have such radiatively-inefficient cores with, on average, $L_{\text{kin}}/L_{\text{nuc}}$ exceeding 200. Explanations of this puzzle carry significant implications for the origin and operation of jets, as well as on establishing the importance of kinetic feedback for the evolution of galaxies and their surrounding medium.

Keywords. X-rays: galaxies: clusters, (galaxies:) cooling flows, galaxies: jets

1. Introduction

Clusters of galaxies with steeply rising X-ray surface brightness profiles are known as cool core clusters, and require strong feedback from their central AGN to offset cooling of the intracluster medium (ICM).

For highly-luminous cool core clusters (with $L_X \geq 10^{45}$ erg s⁻¹), the central AGN must be injecting on average 10^{45} erg s⁻¹ into the surrounding medium. We do not know the black hole (BH) mass for most of these AGN but we expect that they lie between 10^9 and $10^{10} M_\odot$, based on the few reliable measurements in the literature (e. g. M87, Macchetto *et al.* 1997). Given that these objects are in an exceptional environment and are active continuously, it is not clear that the standard $M_{\text{BH}} - \sigma$ or $M_{\text{BH}} - M_K$ relations are relevant for them. For $M_{\text{BH}} \sim 10^9 M_\odot$, this means that these black holes must be operating at high enough Eddington rates that they should be radiatively efficient (see Fig. 1). We would therefore expect to see an X-ray point source.

We present a sample of strong cool core, highly-luminous clusters, for which there is *no* evidence of a nuclear X-ray point source in the *Chandra* images. Using these images, we derive upper limits of the nuclear luminosities with the web interface of PIMMS (Mukai 1993), which converts a count rate into an expected flux. We also investigate whether there is a hidden power law in the X-ray spectra, but find no such evidence in any of our objects. Finally, we calculate the energy that must be injected by the AGN in order to offset the cooling (L_{outflow}) of the ICM, and compare it with the nuclear luminosity of the AGN.

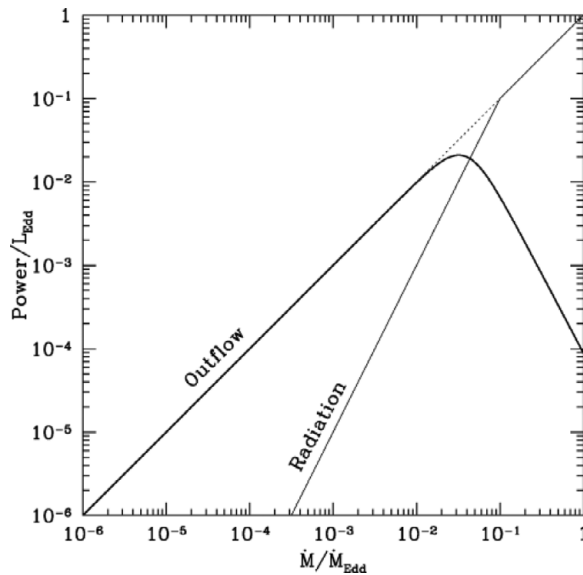


Figure 1. Sketch of black hole energy release as a function of mass accretion (Churazov *et al.* 2005). The energy in radiation dominates at high accretion rates. If a black hole is releasing 10^{45} erg s $^{-1}$, then for $M_{\text{BH}} \sim 10^9 M_{\odot}$ the power exceeds $10^{-2} L_{\text{Edd}}$, indicating that $P_{\text{radiation}} \geq P_{\text{outflow}}$, i.e. we should see a nuclear point source. For $M_{\text{BH}} \geq 10^{10} M_{\odot}$, the power exceeds $10^{-3} L_{\text{Edd}}$. Here, P_{outflow} could dominate over $P_{\text{radiation}}$, if the black hole is in a low accretion state.

2. Results

Our results are shown in Fig. 2 and Fig. 3, and reveal a significant population of objects requiring high kinetic input from an AGN to offset cooling and/or high jet power, yet are without a detected X-ray nucleus. These objects appear to be radiatively inefficient with on average, $L_{\text{kin}}/L_{\text{nuc}} \geq 200$. We examine 7 possible explanations as to why these objects appear to be so radiatively-inefficient.

- First, they could simply be strongly Doppler-suppressed. However, all jets would need to be aligned with the plane of the sky, and it is unlikely that all of our objects have jets with a preferred geometry.

- Second, we could be dealing with Advection Dominated Accretion flows (Narayan & McClintock 2008), but it would be difficult to obtain jet powers of 10^{45} erg s $^{-1}$.

- Next, our objects could have magnetically-dominated black holes, but the mechanism responsible for creating jets in magnetically-dominated accretion discs still remains poorly understood, which makes it difficult to give any definitive conclusion.

- They could also be highly-absorbed. The power must then emerge at longer wavelengths. However, only three of our central AGN have large IR luminosities. Furthermore, this luminosity is from *Spitzer* data, and remains un-resolved for our objects. Most of the far IR emission is also likely due to young stars, since there is much evidence from excess blue light, that there is ongoing star formation in our objects (O’Dea *et al.* 2008).

- We could be looking at the variability of how many X-rays are produced by the core (i.e. the radiative efficiency of the jets). Perhaps the radiative efficiency of a jet depends on how much interaction occurs between the jet and surrounding gas. Extremely powerful steady jets may create a channel which results in little such interaction.

- Next, they could be spin-powered black holes which could allow a strong jet to be created while keeping the AGN radiatively dim (see McNamara *et al.* 2009). A high

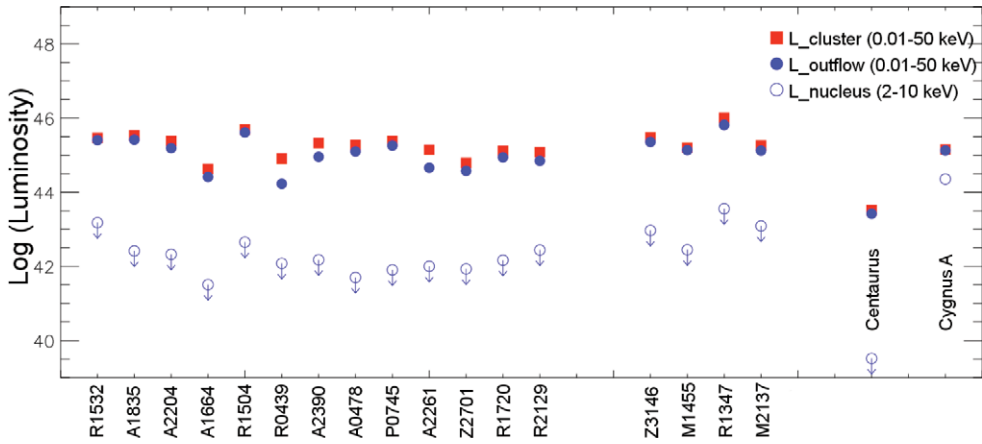


Figure 2. Shown here, is our sample of 13 clusters with no clear X-ray point source. We also include 4 clusters with a possible point source (Z3246, M1455, R1347 and M2137); the Centaurus cluster, known to have a radiatively-inefficient AGN; and Cygnus A, a perfect counter-example with an obvious point source. For each object, we show the total X-ray luminosity ($L_{cluster}$), the energy needed by the AGN to offset cooling ($L_{outflow}$) and the derived upper limit for the nuclear X-ray luminosity ($L_{nucleus}$).

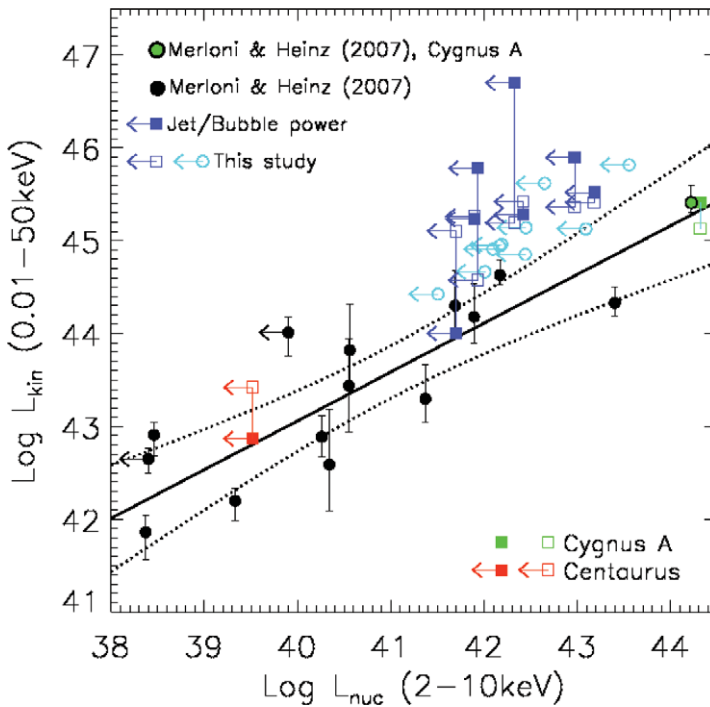


Figure 3. Kinetic bolometric luminosity versus nuclear luminosity in the 2 – 10 keV band. The AGN of Merloni & Heinz (2007) are shown with black filled circles. These authors found that there was a statistical correlation between the kinetic bolometric luminosity and nuclear luminosity of an AGN. Their sample included 15 AGN, and we plot with a black line, the linear fit, and with the dotted lines, the 2σ limits. Our results are shown with the non-filled symbols. Clusters with estimates of their kinetic luminosity from measurements of the energy output from bubbles/jets are shown in squares and the estimate is shown with the filled square.

spin parameter would then be required, which according to our significant population of objects, would mean that highly rotating BHs are not rare.

- Finally, a promising explanation could be that our objects have an ultra-massive black hole ($M_{\text{BH}} \geq 10^{10} M_{\odot}$), then the radiative power could fall well below the kinetic power as shown in Fig. 1.

3. Summary

We have identified a population of objects which require powerful jets to be present but have no X-ray detectable nucleus. We have also identified a range of possible explanations, each of which carries significant implication for the origin and operation of jets. The black holes may be ultramassive ($M_{\text{BH}} \gg 10^9 M_{\odot}$), or have very high spin, or be highly obscured. They may also be mostly off, yet unobservable when they are switched on, or have highly radiatively inefficient jets.

References

- Churazov, E., Sazonov, S., Sunyaev, R., Forman, W., Jones, C., & Böhringer, H. 2005, *MNRAS* (Letters), 363, L91
- *Hlavacek-Larrondo, J. & Fabian, A. C. 2010, accepted to *MNRAS*, arXiv: 1007.1974
- Macchetto, F., Marconi, A., Axon, D. J., Capetti, A., Sparks, W., & Crane, P. 1997, *ApJ*, 489, 579
- McNamara, B. R., Kazemzadeh, F., Rafferty, D. A., Birzan, L., Nulsen, P. E. J., Kirkpatrick, C. C., & Wise, M. W. 2009, *ApJ*, 698, 594
- Merloni, A. & Heinz, S. 2007, *MNRAS*, 381, 589
- Mukai, K. 1993, *Legacy*, 3, 21
- Narayan, R. & McClintock, J. E. 2008, *New Astron. Revs*, 51, 733
- O’Dea, C. P., Baum, S. A., Privon, G., Noel-Storr, J., Quillen, A. C., Zufelt, N., Park, J., Edge, A., Russell, H., Fabian, A. C., Donahue, M., Sarazin, C. L., McNamara, B., Bregman, J. N., & Egami, E. 2008, *ApJ*, 681, 1035