

6. STELLAR PHOTOMETRY WITH MODERN ARRAY DETECTORS

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INTRODUCTION AND BASIC REFERENCES FOR STELLAR PHOTOMETRY WITH CCD

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1 Project and aims of the meeting

An object of this meeting is to define and emphasize the best practices that allow stellar photometry done with array detectors to be as accurate as possible. We dream of applying simple and clearly justified principles; in reality, we find the equipment has physical and technical properties that interfere with each other. Photometry is a method, a metrology; so we first have to define the desired accuracy. Some will be happy with 0.1 magnitude; others need 0.01, or even 0.001 magnitude. The necessary precautions must be matched to these ambitions.

2 Conditions for stellar photometry

The first requirement is precision of measurements. Although there are some difficulties, there are already means of extracting the signal with high precision. This topic has been studied most thoroughly, but some difficulties remain, particularly when fields are crowded.

The second requirement is for known, stable passbands, which must match the standard ones. This point has frequently been underestimated. If we all want to measure the same thing, accurately as well as precisely, it is of first priority to fix and check the right passband. Later transformations by mathematical recipes cannot recover the physical information that has not been included in the measurements. When using CCD's, it is vital to take the greatest possible care about this point.

The reduction to outside the atmosphere, and the relation to a standard system is also a problem. It is necessary to observe enough standard stars, of the right distribution of colors and magnitudes. To have adequate sets of such stars remains a challenge.

The above, unavoidable requirements cannot easily and perfectly be fulfilled with an array detector.

However, if the following conditions are met:

- a) The matching of the passbands is tested and good enough.
- b) The knowledge and control of the stability of all the functions of the camera (filters and detector) are ascertained.

- c) The chromatic properties of the instrument (the colour equation) are methodically and periodically evaluated relatively to a standard having an extended range in colours and magnitudes.
- d) The star's signals are precisely deconvolved and summed-up.
- e) Some stars within each field are well known in the given system.

It should then be possible to reach or even improve the threshold of accuracy required.

3 Basic references for stellar photometry with CCD

In the following I have attempted to set up, with the help of several experts in the field, a number of references, of course non-exhaustive, which are listed as follows:

3.1 General overview

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3.3 Software packages and descriptions

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3.4 Specific problems

3.4.1 PSF fitting

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