

Far Infrared Luminosity Function of Local Galaxies in the AKARI Deep Field South

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Abstract. We present the first far-infrared luminosity function in the AKARI Deep Field South, a premier deep field of the AKARI Space Telescope, using spectroscopic redshifts obtained with AAOmega. To date, we have found spectroscopic redshifts for 389 galaxies in this field and have measured the local ($z < 0.25$) $90 \mu\text{m}$ luminosity function using about one-third of these redshifts. The results are in reasonable agreement with recent theoretical predictions.

Keywords. galaxies: evolution, galaxies: luminosity function, infrared: galaxies

1. AKARI Deep Field South

The AKARI Deep Field South (ADF-S) is centered on RA 4h 44m 00s, Dec $-53^\circ 20' 00''$ (J2000), and is a low-cirrus ($I_{100\mu\text{m}} < 0.5 \text{ MJy sr}^{-1}$) region of about 12 square degrees near the South Ecliptic Pole (Matsuhara *et al.* 2006). It has been extensively studied over the full area by the AKARI-FIS (at 60, 90, 140, and 160 μm , Shirahata *et al.* 2008), Spitzer (24 and 70 μm , Clements *et al.* 2011), BLAST (250, 350, 500 μm , Valiante *et al.* 2010), and over part of the area by CTIO (UVBRI), AKARI-IRC (7 mid-infrared bands, Pearson *et al.* in preparation), ATCA (deep radio survey at 1.4 GHz, White *et al.* 2011), ASTE/AzTEC (1.1 mm, Hatsukade *et al.* 2011) and APEX/LABOCA (870 μm , Khan *et al.* in preparation). Other work includes CIB detection (Matsuura *et al.* 2011) and cross-identification with public data and SEDs (Malek *et al.* 2010). A large part of the field is being mapped by Herschel as part of the HerMES project.

The 90 μm catalogue from the AKARI Far Infrared Surveyor (FIS) deep survey lists 2,282 sources at a signal-to-noise ratio $\text{SNR} > 5$ (giving a minimum flux of 12.81 mJy). Since this wavelength is within the range at which the re-processed radiation received from dust in Star Forming Galaxies (SFGs) is expected to peak, the evolution of the 90 μm luminosity function provides an excellent proxy for the study of star formation.

This poster (which is based on Sedgwick *et al.* 2011) describes the first luminosity function from this data, for $z < 0.25$ galaxies, and combines this data with earlier ISO data from the ELAIS survey to provide the highest signal-to-noise deep luminosity function at this wavelength to date. We also compare our results with predictions from a recent backward galaxy evolution model, finding a reasonable agreement. Our aim is to provide

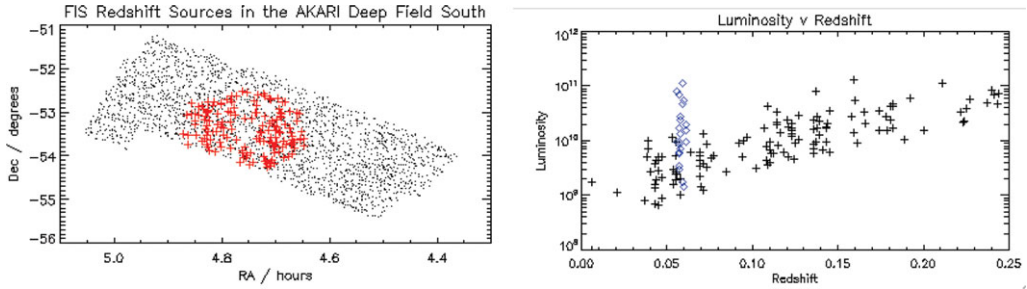


Figure 1. Left: FIS sources with $z < 0.25$ (plus symbols) on all FIS sources (dots) in the AKARI Deep Field South. Right: the redshift - luminosity plane for all sources used, with non-cluster sources shown as plus symbols and the cluster sources shown as diamonds. Luminosity is $\nu L_\nu / L_\odot$ as calculated for the luminosity functions in this paper.

a local benchmark against which to compare luminosity functions of more distant SFGs based on more detailed, deeper data which will become available in the near future.

2. AAOmega Spectroscopy

We have used AAOmega to obtain spectra of selected infrared and sub-millimetre sources with optical counterparts in 3.14 square degrees in the centre of the ADF-S. Data was obtained between October 2007 and November 2008 and redshifts identified for 389 sources, of which we have used 130 for our luminosity function.

Redshifts were measured using the $[\text{NII}]/\text{H}\alpha/[\text{NII}]$, and $\text{H}\beta/[\text{OIII}]/[\text{OIII}]$ emission lines. Redshifts were only taken where one or other (or both) of these two triplets of emission lines were identified. Unless out of range, the $[\text{SII}]$ doublet was also seen. The $[\text{OII}]$ emission line resulting from the $\lambda\lambda 3726, 3729$ was often found at the same redshift to support the identifications. AGNs and SFGs were segregated using a BPT diagram. Figure 1 (left) shows the redshifts $z < 0.25$ which were used for this paper.

As shown in the luminosity/redshift plot (Figure 1 right), we have identified a new cluster in this field at $z \sim 0.06$, centred on RA 4h 42m 12.5s, Dec $-53^\circ 30' 46''$ which is discussed in Sedgwick *et al.* (2011) and excluded from the luminosity function.

3. Far-infrared Local Luminosity Function

In the absence of completeness as a function of flux based directly on AKARI data (under construction, Shirahata *et al.*), we have used the ultra-deep Spitzer $70 \mu\text{m}$ observations of the GOODS-N field which reached down to a flux of 1.2 mJy. Frayer *et al.* (2006) showed that this is well fitted by the model of Lagache *et al.* (2004), which we used to generate number counts against which to measure AKARI completeness. We have adjusted for our $90 \mu\text{m}$ case using the model described in Pearson (2001), scaling the fluxes by a factor of 0.673 on the basis of its prediction that $N(S_{70\mu\text{m}} > 0.0673 \text{ Jy}) = N(S_{90\mu\text{m}} > 0.1 \text{ Jy})$. In addition, we have estimated spectroscopic completeness by comparing the magnitude distribution of AAOmega redshift sources with that of APM B-magnitude sources which have counterparts in the total ADFS-FIS catalogue. The area of the sample was re-normalised to account for large scale cosmic variance, using the number density of objects $> 0.1 \text{ Jy}$ in the 50 deg^2 SWIRE survey (following the procedure in ELAIS, Serjeant *et al.* 2001). We calculated the $1/V_{\text{max}}$ luminosity function (Schmidt 1968) following the methodology in Serjeant *et al.* (2004). K-corrections were made assuming the M82 star-forming spectral energy distribution. We have assumed cosmological parameter values of $H_0 = 72 \text{ kms}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

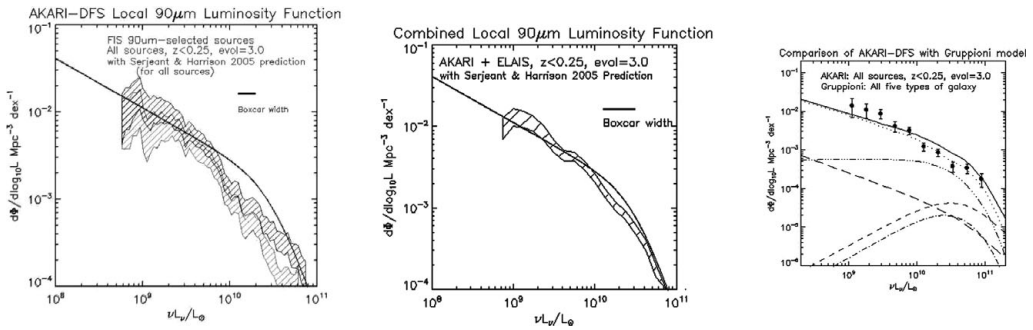


Figure 2. Left: Our new ADF-S luminosity function (Poisson error bands hatched; SFGs-only in shadow). The line is the double power-law prediction by Serjeant & Harrison 2005. Centre: Combined AKARI and ELAIS luminosity function. Right: Comparison of our (unsmoothed) data with predictions from the Gruppioni *et al.* (2011) model. Lines (from top at 10^{10}): total (solid), spirals (dotted), Seyfert2s (dashed-three dots), AGN1s (long-dashed), starbursts (short-dashed), obscured AGNs (dashed-dotted).

In Figure 2, the left panel shows our luminosity function assuming pure luminosity evolution of $(1+z)^3$. We have combined our results with the $90 \mu\text{m}$ luminosity function of the ELAIS-N fields presented in Serjeant *et al.* (2004) in the centre panel (as described in detail in Sedgwick *et al.* 2011), and in the right panel we show that our results are in reasonable agreement with the theoretical prediction for the $90 \mu\text{m}$ luminosity function for $0 < z < 0.25$ galaxies from a new backward evolution model of Gruppioni *et al.* (2011) which includes separate evolution for five types of galaxy. The figure also shows that our result is close to the Serjeant & Harrison (2005) double power-law prediction, with a slight faint-end excess and a deficit just below L^* , attributed to large-scale structure.

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