

## Research Paper

# The GLEAM 200-MHz local radio luminosity function for AGN and star-forming galaxies

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## Abstract

The Galactic and Extragalactic All-sky Murchison Widefield Array (GLEAM) is a radio continuum survey at 76–227 MHz of the entire southern sky (Declination  $< +30^\circ$ ) with an angular resolution of  $\approx 2$  arcmin. In this paper, we combine GLEAM data with optical spectroscopy from the 6dF Galaxy Survey to construct a sample of 1 590 local (median  $z \approx 0.064$ ) radio sources with  $S_{200 \text{ MHz}} > 55 \text{ mJy}$  across an area of  $\approx 16\,700 \text{ deg}^2$ . From the optical spectra, we identify the dominant physical process responsible for the radio emission from each galaxy: 73% are fuelled by an active galactic nucleus (AGN) and 27% by star formation. We present the local radio luminosity function for AGN and star-forming (SF) galaxies at 200 MHz and characterise the typical radio spectra of these two populations between 76 MHz and  $\sim 1$  GHz. For the AGN, the median spectral index between 200 MHz and  $\sim 1$  GHz,  $\alpha_{\text{high}}$ , is  $-0.600 \pm 0.010$  (where  $S \propto \nu^\alpha$ ) and the median spectral index within the GLEAM band,  $\alpha_{\text{low}}$ , is  $-0.704 \pm 0.011$ . For the SF galaxies, the median value of  $\alpha_{\text{high}}$  is  $-0.650 \pm 0.010$  and the median value of  $\alpha_{\text{low}}$  is  $-0.596 \pm 0.015$ . Among the AGN population, flat-spectrum sources are more common at lower radio luminosity, suggesting the existence of a significant population of weak radio AGN that remain core-dominated even at low frequencies. However, around 4% of local radio AGN have ultra-steep radio spectra at low frequencies ( $\alpha_{\text{low}} < -1.2$ ). These ultra-steep-spectrum sources span a wide range in radio luminosity, and further work is needed to clarify their nature.

**Keywords:** radio continuum: galaxies — surveys — methods: data analysis — catalogues — galaxies: active — galaxies: starburst

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## 1. Introduction

The local radio luminosity function (RLF) is a measure of the variation in the global average space density of radio sources with luminosity at the present epoch. It provides an essential benchmark from which to analyse the cosmic evolution of active galactic nuclei (AGN) and star-forming (SF) galaxies to high redshift (see e.g. Condon, 1989) and is a key observable in extragalactic radio source population models.

The radio source population in the local universe has to date been best studied at 1.4 GHz by combining large-area radio continuum and optical redshift surveys (Condon, Cotton, & Broderick, 2002; Sadler et al., 2002; Mauch & Sadler, 2007; Best & Heckman, 2012). Members of the two source classes (AGN and SF) can usually be distinguished using optical spectra. The local RLFs of these two source populations cross over at a 1.4-GHz radio luminosity of  $\approx 10^{23} \text{ W Hz}^{-1}$ , SF galaxies dominating the population of radio sources below this power and radio-loud AGN above it.

In constraining radio source population models, it is important to use samples selected at a wide range of frequencies because of the changing nature of the sources contributing to the local RLF with frequency.

The radio emission from radio-loud AGN at low frequencies ( $\lesssim 200 \text{ MHz}$ ) mainly arises from the radio lobes rather than

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**Table 1.** Comparison of radio-selected samples used to derive the local RLF at frequencies between 150 MHz and 20 GHz. The samples extend to a maximum redshift of 0.2–0.3, corresponding to a look-back time of  $\approx 2.4 - 3.4$  Gyr.

Sample	Magnitude and flux density limits	Area (deg <sup>2</sup> )	AGN		SF galaxy	
			Number	Median $z$	Number	Median $z$
GLEAM-6dFGS (this paper)	$K < 12.65, S_{200 \text{ MHz}} > 55 \text{ mJy}^a$	16 679 <sup>a</sup>	1 157	0.081	427	0.015
LoTSS DR1-SDSS (Sabater et al. 2019)	$14.5 < r < 17.77, S_{150 \text{ MHz}} \gtrsim 0.5 \text{ mJy}$	424	2 121	0.143	8 494	0.097
NVSS-2MASX (Condon, Matthews, & Broderick, 2019)	$K \leq 11.75, S_{1.4 \text{ GHz}} \geq 2.45 \text{ mJy}$	23 032	2 763	0.12	6 699	0.06
NVSS/FIRST-GAMA (Pracy et al. 2016)	$m_i < 20.5, S_{1.4 \text{ GHz}} > 2.8 \text{ mJy}$	$\approx 900$	1 692	0.15–0.20 <sup>b</sup>	527	0.05–0.10 <sup>b</sup>
GAMA-GMRT (Prescott et al. 2016)	$r \leq 19.8, S_{325 \text{ MHz}} \gtrsim 5 \text{ mJy}$	138	134	0.14–0.16 <sup>b</sup>	38	0.06
AT20G-6dFGS (Sadler et al. 2014)	$K < 12.75, S_{20 \text{ GHz}} > 50 \text{ mJy}$	16 980	202	0.058	–	–
NVSS-SDSS (Best & Heckman 2012)	$14.5 < r < 17.77, S_{1.4 \text{ GHz}} > 5 \text{ mJy}$	$\approx 7 100$	7 302	0.16	1 866	0.055
NVSS-6dFGS (Mauch & Sadler 2007)	$K < 12.75, S_{1.4 \text{ GHz}} > 2.8 \text{ mJy}$	7 076	2 661	0.073	4 006	0.035

<sup>a</sup> For the GLEAM-6dFGS sample, there are two regions of different depths. The flux density limit in the deep region, and the total area of the deep and shallow regions are quoted in the table. The flux density limit in the shallow region is 100 mJy. The area of the deep region is 5 113 deg<sup>2</sup> and the area of the shallow region is 11 566 deg<sup>2</sup>.

<sup>b</sup> For the NVSS/FIRST-GAMA and GAMA-GMRT samples, the median redshifts are estimated from the redshift histograms.

the radio core, hotspots, and jets that dominate the emission of sources at high radio frequencies. Since the radio lobes are not subject to relativistic beaming or ‘Doppler boosting’ (see e.g. Blandford & Königl, 1979), low-frequency surveys allow radio-loud AGN to be selected independently of the orientation of the jet axis. Low-frequency surveys can also reveal past activity that is not evident at higher frequencies: large-scale, low-frequency radio emission evolves on relatively long timescales, providing a measure of the long-term jet activity. High-frequency radio emission generally evolves on much shorter timescales, providing a better measure of current jet activity (see e.g. Sadler et al., 2006; Hurley-Walker et al., 2015).

Table 1 compares some of the most recent radio-optical samples used to derive the local RLF. At high frequencies, Sadler et al. (2014) measured the local RLF at 20 GHz by matching radio sources from the Australia Telescope 20 GHz (AT20G) survey (Murphy et al., 2010) with nearby galaxies from the Third Data Release of the 6dF Galaxy Survey (6dFGS DR3; Jones et al., 2009). Although their sample contained some FRI and FRII radio galaxies, it was dominated by compact (FR-0; Ghisellini, 2011) radio AGN without any extended radio emission apparent at lower frequencies. The observed properties of these compact 20 GHz sources are consistent with them being a mixed population including young Compact Steep-Spectrum (CSS) and Gigahertz-Peaked Spectrum (GPS) radio galaxies (O’Dea, 1998). The AT20G catalogue included very few galaxies where the radio emission arose mainly from star formation processes, so the 20-GHz RLF was only measured for the AGN population.

At low frequencies, Prescott et al. (2016) used data from the Giant Metrewave Radio Telescope (GMRT) to measure the 325-MHz local RLF for a small sample of nearby ( $z < 0.25$ ) radio-detected AGN and SF galaxies in a 138 deg<sup>2</sup> area covered by the Galaxy and Mass Assembly (GAMA; Driver et al., 2009) survey. More recently, Sabater et al. (2019) derived the local ( $z < 0.3$ ) RLF at 150 MHz for radio AGN and SF galaxies separately for a much larger sample by cross-matching the first data release of the LOFAR Two-metre Sky Survey (LoTSS DR1; Shimwell et al., 2019) with the Sloan Digital Sky Survey (SDSS; York et al., 2000; Stoughton et al., 2002) main galaxy spectroscopic sample. Their radio AGN luminosity function is in good agreement with previous determinations of the local RLF at 1.4 GHz assuming a spectral index  $\alpha = -0.7$  (where  $S \propto \nu^\alpha$ ).

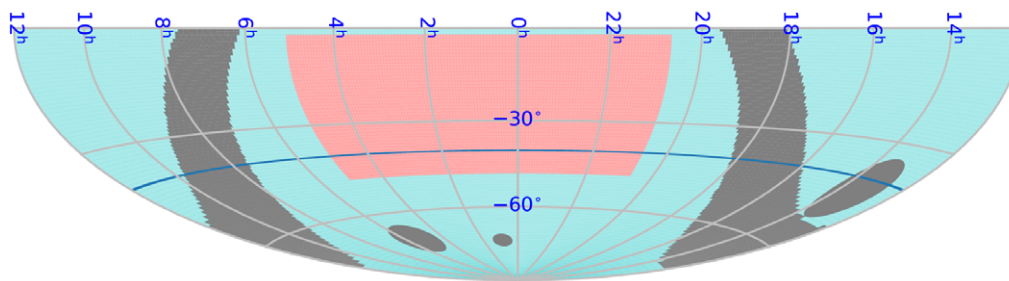
In this paper, we study the local radio source population at 200 MHz over most of the southern sky (16 679 deg<sup>2</sup>) by combining data from the GaLactic and Extragalactic All-sky Murchison Widefield Array (GLEAM) survey (Wayth et al., 2015) with 6dFGS DR3. GLEAM is a wide-area radio continuum survey at 76–227 MHz with an angular resolution of  $\approx 2$  arcmin, conducted with the Murchison Widefield Array (MWA; Tingay et al., 2013). The large fractional bandwidth of the MWA makes it possible to measure in-band spectral indices for the vast majority of the sources detected in GLEAM. In addition, the compact antenna layout of the MWA gives it extremely high surface brightness sensitivity, which is important for recovering extended radio emission in nearby galaxies.

We measure the local RLF for AGN and SF galaxies and characterise the typical radio spectra of these two populations. Our local radio sample is far shallower than the LoTSS-SDSS sample by Sabater et al. (2019) but covers a far wider area of sky, has a lower median redshift, and contains a larger number of AGN at high radio luminosities ( $P_{200 \text{ MHz}} > 10^{24} \text{ W Hz}^{-1}$ ). All sources in the sample have measured spectral indices between 200 MHz and  $\sim 1$  GHz obtained using the NRAO VLA Sky Survey (NVSS; Condon et al., 1998) at 1.4 GHz and the Sydney University Molonglo Sky Survey (SUMSS; Mauch et al., 2003) at 843 MHz. In addition, the vast majority of the sources have GLEAM intra-band spectral indices, allowing their radio spectral properties to be studied in a systematic way and providing important information on the nature of nearby radio galaxies down to low luminosities.

The layout of the paper is as follows. The data used in this work are summarised in Section 2. We define the GLEAM-6dFGS sample in Section 3 and describe the radio properties of the sample in Section 4. In Section 5, the local RLF for AGN and SF galaxies is presented and compared with other measurements. Our results are summarised in Section 6. Throughout this paper, we assume a Hubble constant of  $70 \text{ km s}^{-1} \text{ Mpc}^{-1}$  ( $h = 0.70$ ), and matter and cosmological constant density parameters of  $\Omega_M = 0.3$  and  $\Omega_\Lambda = 0.7$ . Right ascension is abbreviated as RA and declination is abbreviated as Dec.

## 2. Data

We describe the radio and optical data used to define the GLEAM-6dFGS sample.



**Figure 1.** The deep (both years: red) and shallow (year one only: cyan) regions of GLEAM used to define the GLEAM-6dFGS sample. The strip at Galactic latitude  $|b| < 10^\circ$  and a few regions surrounding Centaurus A and the Magellanic Clouds, shown in grey, do not contain any GLEAM-6dFGS sources as they are not covered by GLEAM Exgal. The GLEAM-6dFGS sample is restricted to  $\text{Dec} < 0^\circ$  due to the 6dFGS coverage. Flux densities at 1.4 GHz from NVSS are obtained above  $\text{Dec} -39.5^\circ$  (blue line) and flux densities at 843 MHz from SUMSS are obtained below  $\text{Dec} -39.5^\circ$ . The Aitoff map projection is used.

### 2.1. GLEAM Exgal and SGP catalogues and images (76–227 MHz)

We use the GLEAM Extragalactic (Exgal; Hurley-Walker et al., 2017) and South Galactic Pole (SGP; Franzen et al., 2021) data releases as our low-frequency basis data. GLEAM Exgal is based on the first year (2013–2014) of GLEAM observations. It covers  $24\,831 \text{ deg}^2$  at  $\text{Dec} < +30^\circ$ , excluding the strip at Galactic latitude  $|b| < 10^\circ$  and a few regions around bright, complex sources such as the Magellanic Clouds. The GLEAM Exgal catalogue contains 307 455 source components with 20 separate flux density measurements between 76 and 227 MHz selected from a wide-band image centred at 200 MHz, with an angular resolution of  $\approx 2$  arcmin. The typical rms noise in the wide-band image is  $\approx 10$  mJy/beam. Spectral indices between 76 and 227 MHz were derived by fitting a power law to the sub-band flux densities.

GLEAM SGP is based on a subset of both years (2013–2015) of GLEAM observations. It covers an area of  $5\,113 \text{ deg}^2$  surrounding the South Galactic Pole at  $20^{\text{h}}40^{\text{m}} < \text{RA} < 05^{\text{h}}04^{\text{m}}$  and  $-48^\circ < \text{Dec} < -2^\circ$ . The GLEAM SGP catalogue contains 108 851 source components with 20 separate flux density measurements between 76 and 227 MHz selected from a wide-band image centred at 216 MHz, with an angular resolution of  $\approx 2$  arcmin. The typical rms noise in the wide-band image is  $\approx 5$  mJy/beam. This is still well above the classical confusion limit of  $\approx 2$  mJy/beam (Franzen et al., 2019).

Spectral indices across the GLEAM band were derived by fitting a power law to the sub-band flux densities. The catalogue also contains integrated flux densities at 200 MHz,  $S_{200 \text{ MHz}}$ , as in the GLEAM Exgal catalogue. If the GLEAM intra-band spectrum could be well fitted by a power law,  $S_{200 \text{ MHz}}$  was derived from the power-law fit. Otherwise,  $S_{200 \text{ MHz}}$  was derived by extrapolating the 216-MHz integrated flux density from the wide-band image to 200 MHz assuming  $\alpha = -0.8$ .

We refer to the area of sky covered by the GLEAM SGP catalogue, which is depicted by the red shading in Figure 1, as the ‘deep’ region. We refer to the area of sky covered by the GLEAM Exgal catalogue at  $\text{Dec} < 0^\circ$  and outside the deep region as the ‘shallow’ region. The shallow region is depicted by the cyan shading in Figure 1. The deep and shallow regions are fully described in Table 2. In constructing the GLEAM-6dFGS sample, we use the GLEAM SGP catalogue in the deep region and the GLEAM Exgal catalogue in the shallow region.

We use cutouts from the GLEAM Exgal and SGP wide-band images for visual inspection of the cross-identifications outlined in Section 3.4.

### 2.2. SUMSS catalogue and images (843 MHz)

SUMSS covers the entire extragalactic sky ( $|b| > 10^\circ$ ) south of  $\text{Dec} -30^\circ$  at 843 MHz. It has a limiting peak brightness of 6 mJy/beam at  $\text{Dec} \leq -50^\circ$  and 10 mJy/beam at  $\text{Dec} > -50^\circ$ . The angular resolution is  $45 \times 45 \text{ cosec} |\text{Dec}| \text{ arcsec}^2$ .

We use cutouts from SUMSS for the visual inspection at  $\text{Dec} < -39.5^\circ$ . If necessary, we use the SUMSS catalogue to correct the SUMSS flux density quoted in the NVSS/SUMSS-6dFGS catalogue described in Section 2.4.

### 2.3. NVSS catalogue and images (1.4 GHz)

NVSS covers the entire sky north of  $\text{Dec} -40^\circ$  at 1.4 GHz. It has a  $5\text{-}\sigma$  limit in peak brightness of  $\approx 2.5$  mJy/beam and an angular resolution of 45 arcsec.

We use cutouts from NVSS for the visual inspection at  $\text{Dec} \geq -39.5^\circ$ . If necessary, we use the NVSS catalogue to correct the NVSS flux density quoted in the NVSS/SUMSS-6dFGS catalogue.

### 2.4. The 6dFGS and the NVSS/SUMSS-6dFGS catalogue

#### 2.4.1. The 6dFGS DR3 catalogue

The 6dFGS DR3 catalogue (Jones et al., 2009) contains 125 071 galaxies, making near-complete samples with  $(K, H, J, r_F, b_J) \leq (12.65, 12.95, 13.75, 15.60, \text{ and } 16.75)$ ; the median redshift is 0.053. It covers the whole sky south of  $\text{Dec} 0^\circ$  and outside  $10^\circ$  of the Galactic plane. The redshift completeness of the catalogue is around 85%, with some variation across the survey area.

#### 2.4.2. The NVSS/SUMSS-6dFGS catalogue

As noted by White et al. (2020a), many GLEAM sources have complex structure at low frequencies. Combined with the relatively low spatial resolution of the GLEAM images ( $\sim 2\text{--}3$  arcmin synthesised beam), this makes it challenging to identify the host galaxies of GLEAM sources in an automated way from low-frequency data alone. We, therefore, chose to use higher-frequency radio data from NVSS at 1.4 GHz and SUMSS at 843 MHz, which have higher spatial resolution than GLEAM, as a guide for cross-matching the 6dFGS and GLEAM catalogues as well as a source of radio spectral classifications for the GLEAM-6dFGS radio-matched galaxies.

Mauch & Sadler (2007) used data from NVSS to identify 7 824 radio sources with galaxies brighter than  $K = 12.75$  mag in the Second Incremental Data Release (DR2) of the 6dFGS. These



**Table 2.** Description of the deep and shallow regions of GLEAM used in the analysis of this paper

Description	Region	Area (deg <sup>2</sup> )
<b>Deep region</b>	$-48^\circ < \text{Dec} < -2^\circ$ and $20^{\text{h}}40^{\text{m}} < \text{RA} < 05^{\text{h}}04^{\text{m}}$	<b>5 113</b>
Galactic plane	Absolute Galactic latitude $< 10^\circ$	3 578
Centaurus A	$13^{\text{h}}25^{\text{m}}28^{\text{s}}, -43^\circ 01' 09'', r = 9^\circ$	254
Large Magellanic Cloud	$05^{\text{h}}23^{\text{m}}35^{\text{s}}, -69^\circ 45' 22'', r = 5.5^\circ$	95
Small Magellanic Cloud	$00^{\text{h}}52^{\text{m}}38^{\text{s}}, -72^\circ 48' 01'', r = 2.5^\circ$	20
Peeled sources <sup>†</sup>	Radius of 10 arcmin	$< 1$
<b>Shallow region</b>	Dec $< 0^\circ$ excluding the deep region and the regions listed in the middle rows	<b>11 566</b>

<sup>†</sup> The peeled sources are Hydra A, Pictor A, Hercules A, Virgo A, Tau A, Cygnus A and Cassiopeia A; their positions are listed in Hurley-Walker *et al.* (2017).

authors later used the third and final 6dFGS data release (DR3) to compile a (currently unpublished) catalogue of NVSS-6dFGS and SUMSS-6dFGS cross-identifications at Dec  $\geq -39.5^\circ$  and Dec  $< -39.5^\circ$ , respectively, which was used for this study.

The NVSS/SUMSS-6dFGS catalogue was generated using the same methodology described by Mauch & Sadler (2007) and contains  $\approx 12\,500$  6dFGS radio matches in total (representing around 10% of the 125 071 galaxies in the 6dFGS DR3 catalogue).

A common method of determining the origin (star formation or jets originating from the central supermassive black hole) of the bulk of the radio emission in nearby galaxies is through their optical emission line properties. The optical spectrum of each galaxy in the NVSS/SUMSS-6dFGS catalogue was visually classified by Tom Mauch using the same method applied by Mauch & Sadler (2007). If the spectrum showed evidence of a galaxy with ongoing star formation, with strong narrow emission lines of H $\alpha$  and H $\beta$  typical of HII regions, the source was classified as a SF galaxy. Otherwise, the source was classified as an AGN. The AGN class was further divided into three subclasses: (i) sources with pure absorption line spectra, classed as ‘Aa’; (ii) sources with both absorption lines and weak, narrow emission lines characteristic of Low-Ionisation Nuclear Emission-line Regions (LINERs), classed as ‘Aae’; and (iii) sources with strong nebular emission lines (stronger than any hydrogen Balmer emission lines), classed as ‘Ae’.

It is usually straightforward to tell the dominant physical process responsible for the radio emission from the optical spectra (Sadler *et al.*, 1999; Jackson & Londish, 2000). The only significant complication is with the emission line AGN (Ae and Aae objects) that make up 14% of the GLEAM-6dFGS sample. These are a heterogeneous class of objects where both an AGN and star formation may contribute significantly to the radio emission, the dominant mechanism varying between sources (e.g. Ching *et al.*, 2017). It is likely that some of the weaker radio sources in the Ae/Aae class arise mainly from star formation processes. We return to this point when deriving the local RLF for radio AGN and SF galaxies in Section 5.2.

The NVSS/SUMSS-6dFGS catalogue includes the target name of the object from the 6dFGS data base; the position of the object and its total infrared *K*-band magnitude from the Two Micron All-Sky Survey Extended Source Catalogue (2MASS XSC; Jarrett *et al.*, 2000), as listed in the 6dFGS database; the optical redshift and redshift quality flag, as listed in the 6dFGS catalogue; the spectral classification for galaxies with a good-quality 6dFGS spectrum

from the visual inspection; and the integrated flux density of the radio counterpart in NVSS or SUMSS.

Although NVSS and SUMSS flux densities are included in the NVSS/SUMSS-6dFGS catalogue, all the flux densities used in this paper were checked and re-derived from the original NVSS and SUMSS catalogues.

## 2.5. The AllWISE catalogue in the mid-infrared

The *Wide-field Infrared Survey Explorer* (WISE; Wright *et al.*, 2010) covers the whole sky at 3.4 (W1), 4.6 (W2), 12 (W3), and 22 (W4)  $\mu\text{m}$ . We use the AllWISE data release catalogue by Cutri *et al.* (2013) to characterise the GLEAM-6dFGS sources in the mid-infrared. The angular resolutions are 6.1, 6.4, 6.5, and 12.0 arcsec, and the flux sensitivities at  $5\sigma$  are 0.054, 0.071, 0.73, and 5.0 mJy/beam, in W1, W2, W3, and W4, respectively.

## 2.6. Other radio and optical data

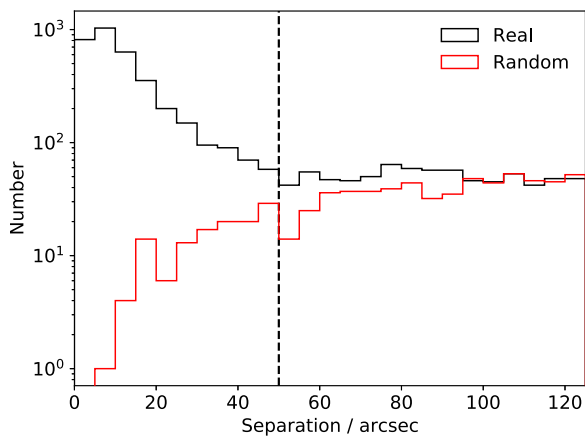
The following data were also used, where available, for the visual inspection to confirm the cross-identification:

- Cutouts from the alternative data release 1 of the TIFR GMRT Sky Survey (TGSS ADR1; Intema *et al.*, 2017). The survey covers the entire sky north of Dec  $-53^\circ$  at 150 MHz to a typical rms sensitivity of 3.5 mJy/beam. The angular resolution is  $25 \times 25$  sec (Dec  $-19^\circ$ ) arcsec<sup>2</sup> at Dec  $< +19^\circ$ .
- Cutouts from the 1.4-GHz Faint Images of the Radio Sky at Twenty Centimetres (FIRST; Becker, White, & Helfand, 1995) survey, which covers  $\approx 10\,575$  deg<sup>2</sup> of the north Galactic cap. The typical sensitivity is 0.13 mJy/beam and the angular resolution 5 arcsec.
- Quick Look images from the first epoch of the VLA Sky Survey (VLASS; Lacy *et al.*, 2020), which covers the entire sky at Dec  $> -40^\circ$  at 2–4 GHz. The sensitivity is  $\approx 120$   $\mu\text{Jy}/\text{beam}$  and the angular resolution 2.5 arcsec.
- Optical blue ( $B_j$ -band) images from the SuperCOSMOS Sky Survey (Hambly *et al.*, 2001) available over the entire southern sky to a typical AB depth of  $B_j < 21$ .

## 3. Definition of the GLEAM-6dFGS sample

### 3.1. Cross-matching GLEAM with 6dFGS, NVSS, and SUMSS

We use the TOPCAT software (Taylor, 2005) to cross-match the NVSS/SUMSS-6dFGS catalogue with the GLEAM SGP catalogue in the deep region and the GLEAM Exgal catalogue in the shallow region. We conduct a simulation to determine the best search radius to automatically accept GLEAM counterparts to 6dFGS galaxies. The black histogram in Figure 2 shows the number of GLEAM sources with identified 6dFGS counterparts as a function of angular separation between the GLEAM and 6dFGS positions. We generate a list of simulated sources by offsetting the GLEAM positions by 10 arcmin in Dec and cross-match the simulated catalogue with the NVSS/SUMSS-6dFGS catalogue. The red histogram shows the number of matches between the simulated catalogue and the NVSS/SUMSS-6dFGS catalogue as a function of angular separation. At a separation of 100 arcsec, the number of false matches is roughly equal to the number of real matches. Matches beyond this separation are therefore unlikely to be genuine. At a separation of 50 arcsec, roughly half the matches are false.



**Figure 2.** Results of a simulation to determine the best search radius to automatically accept GLEAM counterparts to 6dFGS sources. The black histogram shows the number of GLEAM sources with identified 6dFGS counterparts as a function of the angular separation between the GLEAM and 6dFGS positions. The red histogram shows the results obtained for a simulated catalogue generated by offsetting the GLEAM positions by 10 arcmin in Dec. The dashed vertical line at 50 arcsec marks the chosen search radius.

We conclude that a search radius of 50 arcsec provides a good compromise between completeness and reliability. Using this search radius, we obtain a total of 3 495 matches.

### 3.1.1. Completeness and reliability of the cross-matched sample

We use the results of the simulation to estimate the completeness and reliability of the matches. The simulation indicates that 124 random associations with radio sources are expected at separations less than 50 arcsec. This implies that, of the 3 495 accepted matches,  $\approx 124$  are likely to be spurious, corresponding to a reliability of  $1 - \frac{124}{3\,495} = 96.5\%$ . This reliability estimate should be taken as a lower limit because we subsequently remove a number of unreliable matches based on the visual inspection (see Section 3.4.1).

All matches within the 50 arcsec search radius are accepted, hence the completeness is 100% in this zone. We accept no radio source identifications at separations larger than 50 arcsec. However, based on the small excess of real over random matches with position offsets between 50 and 100 arcsec, we expect there to be 176 genuine associations in this zone. This implies that  $\approx 176$  genuine matches are missing from our sample, corresponding to a completeness of  $1 - \frac{176}{3\,495} = 95.0\%$ . The excess of real over random matches at separations greater than 50 arcsec could be, in part, due to clustering of optical galaxies (i.e. these excess matches could be associated with galaxy groups), as noted in previous studies (e.g. Mauch & Sadler, 2007; Best et al., 2005).

### 3.2. Radio completeness

We calculate the flux density levels at which the sample is close to 100% complete in the shallow and deep regions. Hurley-Walker et al. (2017) created maps of the spatial variation of the completeness in GLEAM Exgal at 33 flux density levels spanning the range 25 mJy to 1 Jy. We use these maps to calculate the mean completeness at each flux density level within the shallow region. The mean completeness is found to be  $\approx 90\%$  at  $S_{200\text{ MHz}} = 100$  mJy. The completeness is close to uniform except in a relatively small region of sky covering about  $500\text{ deg}^2$  at  $16^{\text{h}} < \text{RA} < 20^{\text{h}}$  and  $\text{Dec} < -60^\circ$ , where it drops well below 90%.

The GLEAM SGP catalogue is estimated to be 90% complete at  $S_{216\text{ MHz}} = 50$  mJy (Franzen et al., 2021). The completeness is close to uniform across the entire region covered by the catalogue. The completeness limit at 200 MHz will depend on the distribution of spectral indices between 200 and 216 MHz. In estimating the completeness limit at 200 MHz, we take a conservative approach and assume that all sources have a spectral index of  $-1.2$ . Thus, we set the completeness limit at 200 MHz to 55 mJy.

We note that genuine GLEAM-6dFGS matches which are too faint to appear in NVSS/SUMSS will be omitted from our sample. NVSS is close to 100% complete at  $S_{1\,400\text{ MHz}} = 3.4$  mJy. SUMSS has a completeness limit of  $S_{843\text{ MHz}} = 10$  mJy at  $\text{Dec} \leq -50^\circ$  and of  $S_{843\text{ MHz}} = 18$  mJy at  $\text{Dec} > -50^\circ$ . Table 3 shows the limiting spectral index between 200 and 1 400/843 MHz in five different regions of the GLEAM-6dFGS sample depending on the survey data available. The limiting spectral indices assume that the angular size of the source is not significantly larger than 45 arcsec, the angular resolution of NVSS and SUMSS.

The highest limiting spectral index is  $-0.78$  in the section of the deep region covered by SUMSS. The fraction of sources with spectral indices between 200 MHz and  $\sim 1$  GHz,  $\alpha_{\text{high}}$ , steeper than  $-0.78$  is  $\approx 21\%$  based on the distribution of  $\alpha_{\text{high}}$  for the GLEAM-6dFGS sources (see Section 4.3). We therefore expect  $\approx 21\%$  of genuine GLEAM-6dFGS associations to be missing from the sample in this region close to the 55 mJy flux density cut, assuming that the source angular size is  $\lesssim 45$  arcsec. However, this region only represents 5% of the total area covered by the sample.

The limiting spectral indices in the other regions listed in Table 3 lie between  $-1.19$  and  $-1.74$ . We can obtain an estimate of the fraction of sources with  $\alpha < -1.2$  from Seymour et al., in preparation, who cross-matched GLEAM SGP with the third data release from the Australia Telescope Large Area Survey (ATLAS DR3; Franzen et al., 2015); ATLAS DR3 covers the *Chandra* Deep Field South (CDFs; Giacconi et al., 2001), and the European Large Area ISO Survey - South 1 (ELAIS-S1; Oliver et al., 2000) to an rms depth of  $\approx 15\ \mu\text{Jy}/\text{beam}$  at 1.4 GHz. Of the 134 GLEAM SGP sources lying in CDFS and ELAIS-S1, only four (3%) were found to have  $\alpha < -1.2$ . From these spectral index limits, we determine that the GLEAM-6dFGS sample is biased against diffuse radio sources with very steep spectra such as radio haloes and relics. Such sources may be detected by directly cross-matching the GLEAM catalogue with the 6dFGS catalogue, but this is beyond the scope of this paper. We do not expect any bias against this source class to significantly affect the overall completeness of the sample and the measured local RLF.

### 3.3. Source filtering

The determination of the local RLF requires a sample which is complete to the limits of the radio and optical surveys from which it is derived. As described in Section 3.2, the shallow region is complete to  $S_{200\text{ MHz}} = 100$  mJy and the deep region to  $S_{200\text{ MHz}} = 55$  mJy. In order to form a sample which is complete in the radio, we discard all sources with  $S_{200\text{ MHz}} < 100$  mJy and  $S_{200\text{ MHz}} < 55$  mJy in the shallow and deep regions, respectively. Since the 6dFGS DR3 catalogue is close to 100% complete at  $K < 12.65$  mag (Jones et al., 2009), we discard sources with  $K > 12.65$  mag. Finally, we remove sources with unreliable redshifts (redshift quality 1 or 2), those associated with Galactic stars (redshift quality 6) and those with  $z < 0.001$ , leaving a total of 1 688 sources.

**Table 3.** The limiting spectral indices,  $\alpha_{\text{lim}}$ , between 200 and 1400/843 MHz in five different regions of the GLEAM-6dFGS sample depending on the survey data available. Sources with  $\alpha < \alpha_{\text{lim}}$  may not appear in the sample. The final number of GLEAM-6dFGS sources in each region is given in the last column.

Region	Area (deg <sup>2</sup> )	Completeness limit (mJy/beam)				$\alpha_{\text{lim}}$	$N$
		$S_{200\text{ MHz}}$	$S_{843\text{ MHz}}$	$S_{1400\text{ MHz}}$			
GLEAM deep & Dec $\geq -39.5^\circ$	4 340 (26%)	55	–	3.4	–1.43	536	
GLEAM deep & Dec $< -39.5^\circ$	773 (5%)	55	18	–	–0.78	92	
GLEAM shallow & Dec $\geq -39.5^\circ$	6 879 (41%)	100	–	3.4	–1.74	614	
GLEAM shallow & $-39.5^\circ > \text{Dec} \geq -50^\circ$	1 163 (7%)	100	18	–	–1.19	112	
GLEAM shallow & Dec $< -50^\circ$	3 524 (21%)	100	10	–	–1.60	236	

### 3.4. Visual inspection

For all 1 688 selected sources, we visually inspect optical  $J$ -band images of size 15 arcmin from the SuperCOSMOS Sky Survey overlaid with radio contours from GLEAM and NVSS/SUMSS. Firstly, we identify and remove unreliable GLEAM-6dFGS matches. Secondly, we identify which NVSS/SUMSS components are associated with each source and, if necessary, correct the total NVSS/SUMSS flux density. Thirdly, we use the overlays to identify sources whose flux densities in GLEAM are significantly affected by confusion from adjacent sources and apply a deblending method to correct the GLEAM flux densities.

#### 3.4.1. Removing false GLEAM-6dFGS matches

The probability of a GLEAM-6dFGS match being genuine is dependent on the GLEAM-6dFGS offset. The simulations used to determine the best search radius to automatically accept GLEAM-6dFGS identifications (see Section 3.1) indicate that the reliability of the cross-matches is  $\approx 99\%$  for position offsets less than 30 arcsec and  $\approx 73\%$  for position offsets between 30 and 50 arcsec.

We visually inspect the overlays of all sources, paying particular attention to sources with position offsets larger than 30 arcsec and identify 22 false GLEAM-6dFGS matches. These sources are subsequently removed from the sample. The false GLEAM-6dFGS matches are typically the result of confusion with an adjacent source, displacing the position of the GLEAM component away from the 6dFGS position and boosting its flux density; an example of such a case is shown in Figure 3.

#### 3.4.2. Identifying multi-component sources

We attempt to automatically identify multi-component sources in NVSS/SUMSS using the NVSS/SUMSS cutouts. The integration area of a source is taken to consist of all pixels in the NVSS/SUMSS image that are within the contiguous  $3\text{-}\sigma$  contour level, bounding the GLEAM position in question. All catalogued NVSS/SUMSS components located within the integration area are considered to be associated with the GLEAM component. We calculate the brightness-weighted centroid of the NVSS/SUMSS emission from the positions and flux densities of the associated NVSS/SUMSS components. The errors on the positions of the individual NVSS/SUMSS components are assumed to be correlated when calculating the centroid's position error. We also calculate the total flux density of the individual NVSS/SUMSS components. Again, we take a conservative approach and assume that the component flux density errors are correlated.

This automated technique can fail to identify the correct NVSS/SUMSS components in the following two situations: (i) the

NVSS/SUMSS components of a very extended source are well separated and there is no extended emission  $> 3\sigma$  linking the components; (ii) unrelated point sources lie sufficiently close together on the sky such that they are located within the same integration area. In order to identify and correct these errors, we visually inspect all the overlays. Figure 4 shows example overlays for an extended AGN and SF galaxy where the NVSS/SUMSS components are correctly identified using the automated procedure.

Where necessary, we use higher-resolution radio data from FIRST, VLASS, and TGSS to help identify by eye the correct NVSS/SUMSS components. The VLASS Quick Look images can be particularly useful to determine whether overlapping NVSS sources are physically related or not, as shown in Figure 5. In the final GLEAM-6dFGS catalogue, a total of 372 sources (23%) are associated with more than one NVSS/SUMSS component. The source (GLEAM J225615-361754) with the highest number of components is a giant radio galaxy with 10 NVSS components, an angular size of  $\approx 14$  arcmin and a linear size of  $\approx 1.4$  Mpc.

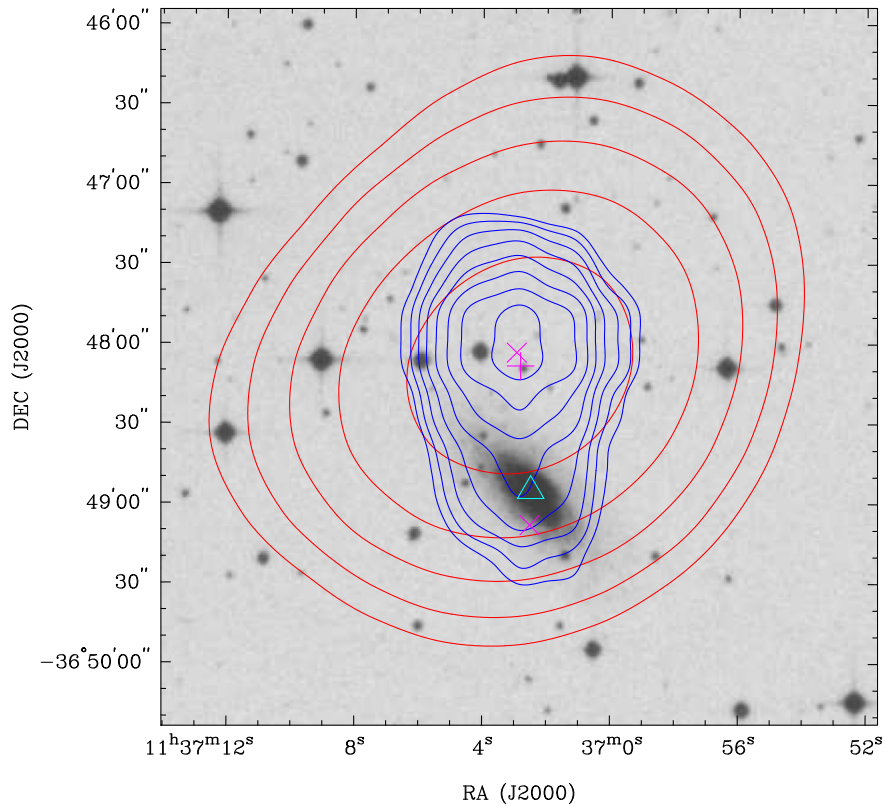
We identify 13 sources that are associated with more than 1 GLEAM component; the highest number of GLEAM components is four. As expected, the number of multi-component sources in GLEAM is much smaller due to the larger GLEAM beam size of  $\approx 2$  arcmin. We sum the integrated flux densities of the individual GLEAM components to obtain the total, integrated flux density at 200 MHz and at each of the 20 sub-band frequencies. When calculating the errors on the total flux densities, we assume that the component flux density errors are correlated. We provide the name(s) of any associated GLEAM component(s) in the GLEAM-6dFGS catalogue.

#### 3.4.3. Correcting the GLEAM flux densities for confusion

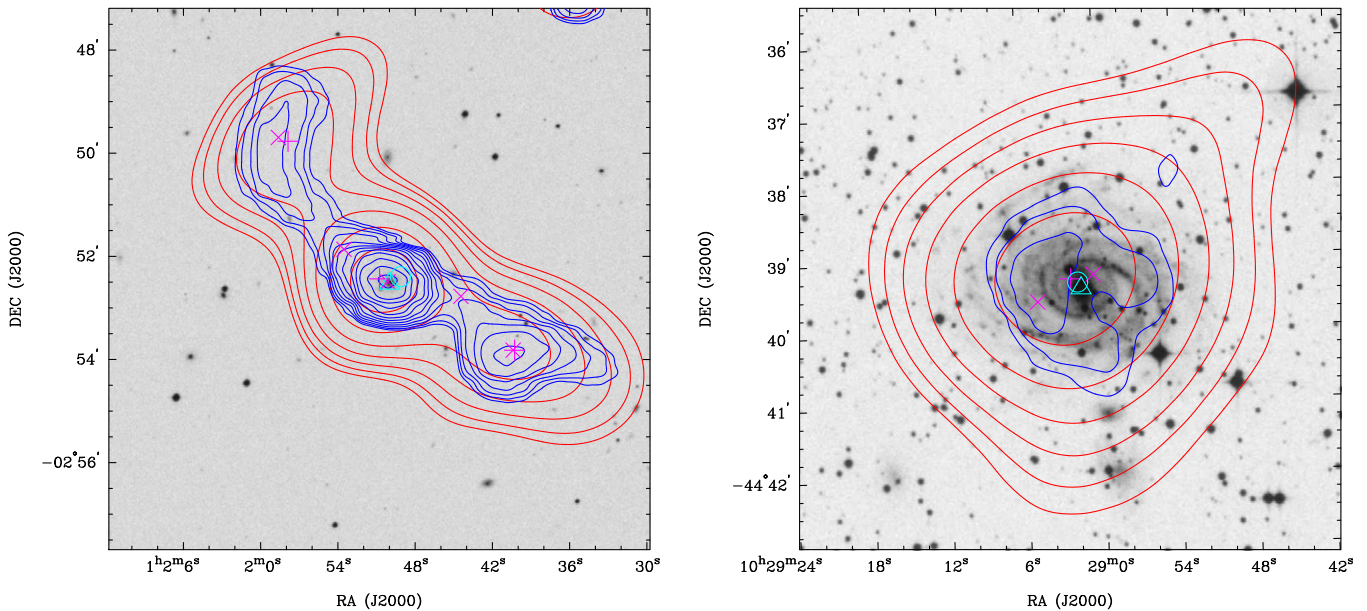
From the visual inspection, we find instances where the MWA beam has blended unrelated NVSS/SUMSS sources together. We identify 254 sources whose GLEAM flux densities are likely to be significantly overestimated due to confusion and that require re-fitting. An example of such a source is shown in Figure 6.

For each of these sources, we use the prioritised fitting mode of the AEGEAN source finder (Hancock *et al.*, 2012; Hancock, Trott, & Hurley-Walker, 2018) to deblend the GLEAM emission from the 6dFGS galaxy. We extract flux densities from the GLEAM wide-band and sub-band images at the positions of all catalogued NVSS/SUMSS components within 10 arcmin from the 6dFGS galaxy, simultaneously. The flux densities of the components are measured using Gaussian fitting. Only the peak flux densities of the components are allowed to vary in the Gaussian fitting; the positions of the components are fixed and their shapes are set to

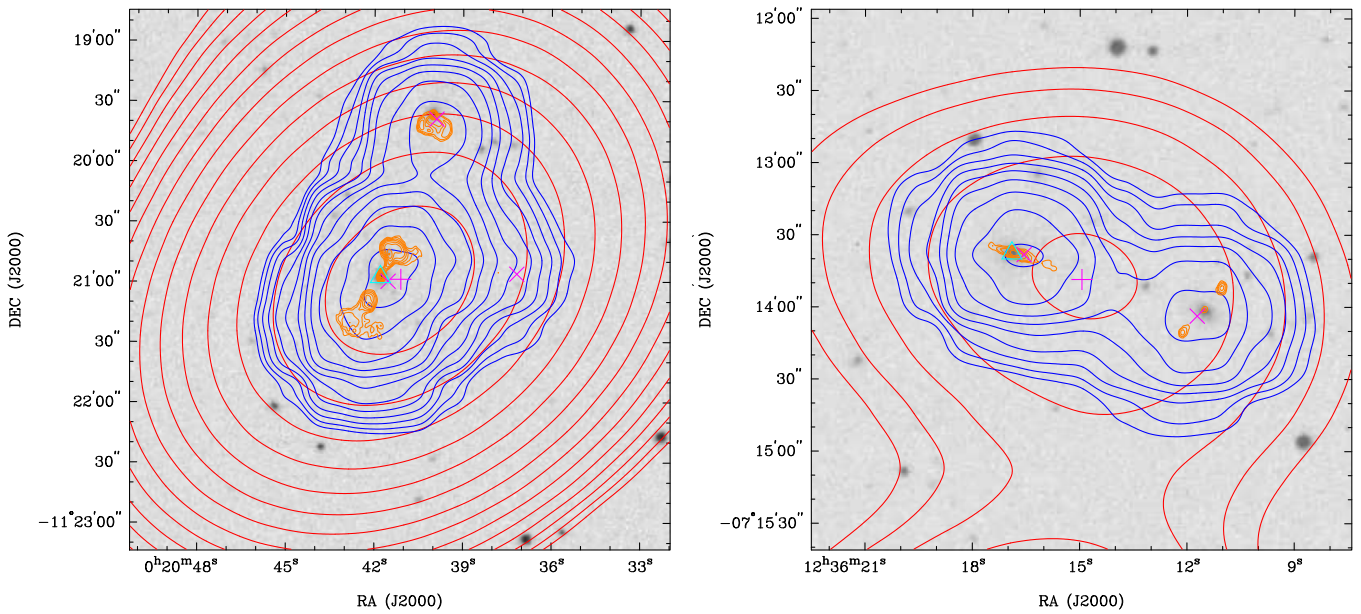




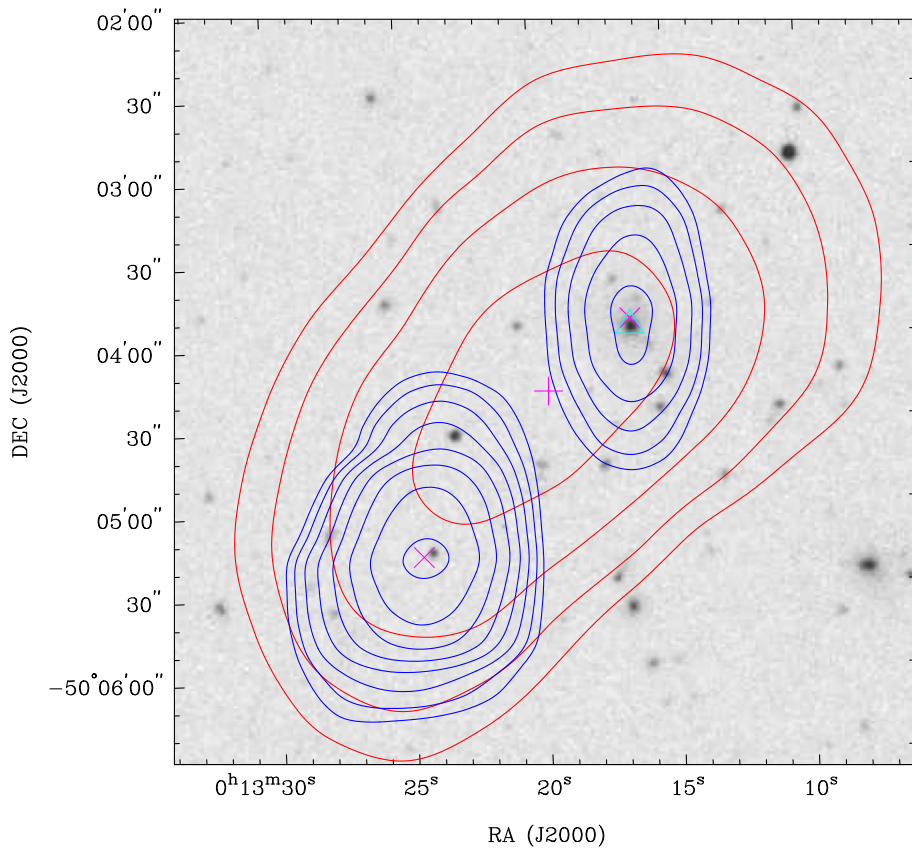
**Figure 3.** Example overlay for a source with a relatively large GLEAM-6dFGS position offset (47.1 arcsec). Radio contours from GLEAM (200 MHz; red) and NVSS (1400 MHz; blue) are overlaid on the SuperCOSMOS *J*-band image (inverted greyscale). For each set of contours, the lowest contour is at the  $3\sigma$  level, where  $\sigma$  is the local rms, with the number of  $\sigma$  increasing by a factor of  $\sqrt{2}$  with each subsequent contour. Catalogue positions from GLEAM (magenta plus signs) and NVSS (magenta crosses) are plotted. The cyan triangle shows the optical position of the source in 6dFGS. The large GLEAM-6dFGS position offset is the result of confusion with an adjacent source. The GLEAM-6dFGS match is not genuine and the source is removed from the sample.



**Figure 4.** Example overlays for an extended FR II at  $z = 0.0706$  (GLEAM J010150-025226; left) and a face-on spiral galaxy at  $z = 0.0084$  (GLEAM J102902-443916; right) in the GLEAM-6dFGS sample. Radio contours from GLEAM (200 MHz; red) and NVSS/SUMSS (1400/843 MHz; blue) are overlaid on the SuperCOSMOS *J*-band image (inverted greyscale). For each set of contours, the lowest contour is at the  $3\sigma$  level, where  $\sigma$  is the local rms, with the number of  $\sigma$  increasing by a factor of  $\sqrt{2}$  with each subsequent contour. Catalogue positions from GLEAM (magenta plus signs) and NVSS or SUMSS (magenta crosses) are plotted. The brightness-weighted centroid position, calculated using the NVSS/SUMSS components, is indicated by a cyan circle. The cyan triangle shows the optical position of the source in 6dFGS.



**Figure 5.** Example overlays including higher-resolution radio data from VLASS to help determine whether overlapping NVSS components are physically related or not. Radio contours from VLASS (2–4 GHz; orange) are plotted, with the lowest contour at the  $5\sigma$  level and the number of  $\sigma$  increasing by a factor of  $\sqrt{2}$  with each subsequent contour. The other contours and symbols are as described in Figure 4. Left: the southern NVSS component is detected as an FRI in VLASS. The northern NVSS component, also detected in VLASS, is physically unrelated. Right: the western NVSS component is detected as a double radio galaxy in VLASS and is physically unrelated with the eastern NVSS component.



**Figure 6.** Example overlay where two distinct galaxies detected in SUMSS are confused in GLEAM. The contours and symbols are as described in Figure 4.



the synthesised beam of the GLEAM image. This is because the Gaussian fitting can be poorly constrained when allowing the positions and shapes of the components to vary. In any case, given the higher resolution of NVSS and SUMSS, the emission in GLEAM can generally be very well modelled with point sources at the positions of the NVSS and SUMSS components.

The deblended flux density of the NVSS/SUMSS component associated with the 6dFGS galaxy is then taken as our new estimate of the GLEAM flux density at each frequency. If more than one NVSS/SUMSS component is associated with the 6dFGS galaxy, we sum the component flux densities to obtain the total, integrated flux density. After applying this procedure, we obtain deblended flux densities at 20 frequencies between 76 and 227 MHz from the sub-band images. We also obtain deblended flux densities at 200 and 216 MHz from the wide-band images in the shallow and deep regions, respectively. The intra-band spectral index is re-measured by fitting a power law to the 20 sub-band flux densities. In the deep region, a best estimate of the 200-MHz deblended flux density is derived using the same method as in the original GLEAM SGP catalogue (see Section 2.1).

For about one-third of the sources that are refitted, the deblended 200-MHz flux density falls below the 200-MHz flux density limit of the GLEAM-6dFGS sample (100 mJy in the shallow region and 55 mJy in the deep region). After removing these sources, the final number of sources in the GLEAM-6dFGS sample is 1 590. In the GLEAM-6dFGS catalogue, we use the ‘refitted flag’ to indicate which sources have refitted GLEAM measurements. We also flag sources for which the GLEAM flux densities, and in some cases also the NVSS/SUMSS flux densities, are judged to be unreliable due to severe confusion (‘confusion flag’) based on the visual inspection. A total of 24 sources are flagged in this way; their flux densities are likely to be overestimated.

### 3.5. Cross-matching with AllWISE

We use the ‘CDS Upload X-Match’ facility of TOPCAT to cross-match the GLEAM-6dFGS sample with the AllWISE catalogue. At least one match within 5 arcsec of the 6dFGS optical position is obtained for 1 575 of the 1 590 GLEAM-6dFGS sources. In the final GLEAM-6dFGS catalogue, we record the name, position, and mid-infrared magnitudes (and errors) of the closest AllWISE match within 5 arcsec of the 6dFGS optical position.

### 3.6. Final catalogue description

In Table 4, we describe the 87 columns of the GLEAM-6dFGS source catalogue. The electronic version of the catalogue is available from Vizier.

We use the GLEAM and NVSS/SUMSS data to calculate two-point spectral indices:

$$\alpha_{\text{high}} = \frac{\ln [S(\nu_1)/S(\nu_2)]}{\beta}, \quad (1)$$

where  $\nu_1 = 200$  MHz,

$$\nu_2 = \begin{cases} 843 \text{ MHz} & \text{if Dec} < -39.5^\circ \\ 1\,400 \text{ MHz} & \text{otherwise,} \end{cases} \quad (2)$$

$S(\nu_1)$  is the integrated flux density at  $\nu_1$  from GLEAM,  $S(\nu_2)$  is the integrated flux density at  $\nu_2$  from NVSS/SUMSS and

$\beta = \ln (\nu_1/\nu_2)$ . The error on  $\alpha_{\text{high}}$  is given by:

$$\Delta\alpha_{\text{high}} = \frac{\sqrt{\left[\frac{\Delta S(\nu_1)}{S(\nu_1)}\right]^2 + \left[\frac{\Delta S(\nu_2)}{S(\nu_2)}\right]^2}}{\beta}. \quad (3)$$

The 200-MHz flux density errors quoted in the GLEAM Exgal and SGP catalogues,  $\sigma_{\text{fit}}(\nu_1)$ , only account for the image noise. We therefore set

$$\Delta S(\nu_1) = \sqrt{\sigma_{\text{fit}}^2(\nu_1) + [\epsilon S(\nu_1)]^2}, \quad (4)$$

where  $\epsilon$  is the GLEAM external flux scale error. The value of  $\epsilon$  is Dec-dependent but at most Decs  $\epsilon = 0.08$  (see Hurley-Walker et al., 2017 and Franzen et al., 2021 for details).

We calculate rest-frame radio luminosities at 200 MHz using a  $k$ -correction based on the spectral index,  $\alpha$ , and assuming the radio emission is synchrotron emission characterised by a power law. The radio luminosity of a source with flux density  $S_\nu$  at redshift  $z$  and luminosity distance  $d_L$  is therefore given by:

$$L_\nu = \frac{4\pi d_L^2 S_\nu}{(1+z)^{\alpha+1}}. \quad (5)$$

In the radio  $k$ -correction, we use the spectral index calculated from the 20 GLEAM sub-band flux densities between 76 and 227 MHz,  $\alpha_{\text{low}}$ , if  $\Delta\alpha_{\text{low}} \leq 0.2$ . This condition is met for 84% of the sources in the catalogue. For the remaining sources with  $\Delta\alpha_{\text{low}} > 0.2$  or no measurement of  $\alpha_{\text{low}}$ , we use  $\alpha_{\text{high}}$  in the radio  $k$ -correction.

We calculate absolute  $K$ -band magnitudes using  $K$ -band  $k$ -corrections derived by Glazebrook et al. (1995) from the evolutionary synthesis models of Bruzual & Charlot (1993) assuming a delta-function burst of star formation at age 5 Gyr. The absolute  $K$ -band magnitude is given by:

$$M_{K,\text{abs}} = M_K - 5 \log_{10} \left( \frac{d_L}{10 \text{ pc}} \right) - K(z), \quad (6)$$

where  $M_K$  is the apparent  $K$ -band magnitude and  $K(z)$  is the  $K$ -band  $k$ -correction given by equation (6) of Glazebrook et al.

White et al. (2020a,b) compiled a complete sample of the ‘brightest’ radio sources ( $S_{151 \text{ MHz}} > 4$  Jy) at Dec  $< 30^\circ$  from the GLEAM Exgal catalogue, the majority of which are AGN with powerful radio jets. The G4Jy sample consists of 1 863 sources and is over 10 times larger than the revised Third Cambridge Catalogue of Radio Sources (3CRR Laing, Riley, & Longair, 1983). A total of 47 G4Jy sources are in common with the GLEAM-6dFGS sample. For reference, the final column of the catalogue (‘G4Jy\_name’) gives the name of the G4Jy source. We note that not all G4Jy sources with 6dFGS hosts are expected to be included in the GLEAM-6dFGS catalogue due to the maximum allowed separation of 50 arcsec between the GLEAM and 6dFGS positions, and the source filtering applied in Section 3.3.

## 4. Radio properties of the GLEAM-6dFGS sample

Table 5 shows the distribution by type of the spectral classification of the GLEAM-6dFGS sample. Of the 1 590 sources, 1 157 (72.8%) are classified as AGN and 427 (26.9%) as SF galaxies. The vast majority of the AGN have pure absorption line spectra typical of giant elliptical galaxies. For six sources (0.4%), the optical spectrum is unclassifiable or unknown; the analysis presented in this section does not include the six sources with ‘unknown’ optical spectra.

**Table 4.** Column numbers, names, units, descriptions and first row entries for the 87 columns in the GLEAM-6dFGS source catalogue. All reported magnitudes are in the Vega system.

Number	Name	Unit	Description	First row entry
1	6dFGS_target_name	–	Target name of the object from the 6dFGS data base	g0000141-251113
2	6dFGS_RAJ2000	°	RA of the object from the 2MASS XSC (J2000)	0.05883
3	6dFGS_DEJ2000	°	Dec of the object from the 2MASS XSC (J2000)	–25.18692
4	z	–	6dFGS measured redshift	0.0852
5	Q	–	6dFGS redshift quality flag as described in Jones <i>et al.</i> (2009)	4
6	K	mag	Total apparent <i>K</i> -band magnitude	11.89
7	K_abs	mag	Total absolute <i>K</i> -band magnitude	–25.85
8	Spectrum_class	–	Classification of the spectrum as defined in Table 5	Aa
9	AllWISE_name	–	Name of the host galaxy in AllWISE	J000014.07-251112.6
10	AllWISE_RAJ2000	°	RA of the host galaxy (J2000)	0.05865
11	AllWISE_DEJ2000	°	Dec of the host galaxy (J2000)	–25.18684
12	W1mag	mag	W1 magnitude of the host galaxy	12.849
13	err_W1mag	mag	Error on W1 magnitude of the host galaxy	0.025
14	W2mag	mag	W2 magnitude of the host galaxy	12.760
15	err_W2mag	mag	Error on W2 magnitude of the host galaxy	0.028
16	W3mag	mag	W3 magnitude of the host galaxy	11.780
17	err_W3mag	mag	Error on W3 magnitude of the host galaxy	0.261
18	W4mag	mag	W4 magnitude of the host galaxy	8.512
19	err_W4mag	mag	Error on W4 magnitude of the host galaxy	–
20	GLEAM_name	hh:mm:ss+dd:mm:ss	Name of the GLEAM component	GLEAM J000012-251112
21	GLEAM_RAJ2000	°	RA of the GLEAM component (J2000)	0.05394
22	GLEAM_DEJ2000	°	Dec of the GLEAM component (J2000)	–25.18671
23	Offset	arcsec	Offset between GLEAM and 6dFGS positions	16.0
24	region	–	Region of the GLEAM survey (deep or shallow)	deep
25	ncmp_GLEAM	–	Number of GLEAM components	1
26	GLEAM_associated_name	–	Name(s) of any associated GLEAM component(s)	–
27	GLEAM_err_abs_flux_pct	–	Percentage error in absolute flux scale	8.0
28	GLEAM_err_fit_flux_pct	–	Percentage error in internal flux scale	2.0
29	S_200	mJy	Integrated flux density at 200 MHz from GLEAM	115.8
30	dS_200	mJy	Error on integrated flux density at 200 MHz from GLEAM	1.6
31	S_076	mJy	Integrated flux density at 76 MHz from GLEAM	246.1
32	dS_076	mJy	Error on integrated flux density at 76 MHz from GLEAM	55.7
33	S_084	mJy	Integrated flux density at 84 MHz from GLEAM	232.2
34	dS_084	mJy	Error on integrated flux density at 84 MHz from GLEAM	42.7
35	S_092	mJy	Integrated flux density at 92 MHz from GLEAM	246.7
36	dS_092	mJy	Error on integrated flux density at 92 MHz from GLEAM	37.5
37	S_099	mJy	Integrated flux density at 99 MHz from GLEAM	243.7
38	dS_099	mJy	Error on integrated flux density at 99 MHz from GLEAM	34.6
39	S_107	mJy	Integrated flux density at 107 MHz from GLEAM	213.8
40	dS_107	mJy	Error on integrated flux density at 107 MHz from GLEAM	21.5
41	S_115	mJy	Integrated flux density at 115 MHz from GLEAM	191.4
42	dS_115	mJy	Error on integrated flux density at 115 MHz from GLEAM	16.3
43	S_122	mJy	Integrated flux density at 122 MHz from GLEAM	186.7
44	dS_122	mJy	Error on integrated flux density at 122 MHz from GLEAM	14.6
45	S_130	mJy	Integrated flux density at 130 MHz from GLEAM	149.0
46	dS_130	mJy	Error on integrated flux density at 130 MHz from GLEAM	12.4
47	S_143	mJy	Integrated flux density at 143 MHz from GLEAM	136.5
48	dS_143	mJy	Error on integrated flux density at 143 MHz from GLEAM	10.3

**Table 4.** Continued

Number	Name	Unit	Description	First row entry
49	S_151	mJy	Integrated flux density at 151 MHz from GLEAM	144.0
50	dS_151	mJy	Error on integrated flux density at 151 MHz from GLEAM	8.6
51	S_158	mJy	Integrated flux density at 158 MHz from GLEAM	142.1
52	dS_158	mJy	Error on integrated flux density at 158 MHz from GLEAM	8.0
53	S_166	mJy	Integrated flux density at 166 MHz from GLEAM	137.0
54	dS_166	mJy	Error on integrated flux density at 166 MHz from GLEAM	7.3
55	S_174	mJy	Integrated flux density at 174 MHz from GLEAM	141.1
56	dS_174	mJy	Error on integrated flux density at 174 MHz from GLEAM	7.2
57	S_181	mJy	Integrated flux density at 181 MHz from GLEAM	126.7
58	dS_181	mJy	Error on integrated flux density at 181 MHz from GLEAM	5.9
59	S_189	mJy	Integrated flux density at 189 MHz from GLEAM	126.2
60	dS_189	mJy	Error on integrated flux density at 189 MHz from GLEAM	6.0
61	S_197	mJy	Integrated flux density at 197 MHz from GLEAM	121.8
62	dS_197	mJy	Error on integrated flux density at 197 MHz from GLEAM	5.6
63	S_204	mJy	Integrated flux density at 204 MHz from GLEAM	113.0
64	dS_204	mJy	Error on integrated flux density at 204 MHz from GLEAM	5.0
65	S_212	mJy	Integrated flux density at 212 MHz from GLEAM	115.0
66	dS_212	mJy	Error on integrated flux density at 212 MHz from GLEAM	4.7
67	S_220	mJy	Integrated flux density at 220 MHz from GLEAM	100.8
68	dS_220	mJy	Error on integrated flux density at 220 MHz from GLEAM	4.3
69	S_227	mJy	Integrated flux density at 227 MHz from GLEAM	101.9
70	dS_227	mJy	Error on integrated flux density at 227 MHz from GLEAM	4.3
71	alpha_low	-	Spectral index between 76 and 227 MHz from GLEAM	-0.86
72	dalpha_low	-	Error on spectral index between 76 and 227 MHz from GLEAM	0.05
73	reduced_chi2_alpha_low	-	Reduced $\chi^2$ value for GLEAM spectral index fit <sup>a</sup>	0.72
74	refitted_flag	-	'1' indicates that refitted GLEAM measurements are used	0
75	confusion_flag	-	'1' indicates severe source confusion in GLEAM	0
76	P_200	W Hz <sup>-1</sup>	Logarithm of the radio luminosity at 200 MHz	24.31
77	ncmp_NVSSorSUMSS	-	Number of NVSS/SUMSS components	1
78	centroid_RAJ2000	°	RA of the centroid position from NVSS/SUMSS (J2000)	0.05754
79	err_centroid_RAJ2000	arcsec	Error on the RA of the centroid position from NVSS/SUMSS	0.81
80	centroid_DEJ2000	°	Dec of the centroid position from NVSS/SUMSS (J2000)	-25.18719
81	err_centroid_DEJ2000	arcsec	Error on the Dec of the centroid position from NVSS/SUMSS	0.90
82	S_NVSSorSUMSS	mJy	Integrated flux density at 1 400/843 MHz from NVSS/SUMSS	28.4
83	dS_NVSSorSUMSS	mJy	Error on integrated flux density at 1 400/843 MHz from NVSS/SUMSS	1.5
84	alpha_high	-	Spectral index between 200 and 1 400/843 MHz	-0.72
85	dalpha_high	-	Error on spectral index between 200 and 1 400/843 MHz	0.05
86	Freq	MHz	Indicates whether NVSS or SUMSS is used	1 400
87	G4Jy_name	-	Name of the source from the G4Jy sample	-

Best & Heckman (2012) proposed a fundamental dichotomy between high-excitation radio galaxies (HERGs), fuelled at high rates through radiatively efficient classic accretion discs, and low-excitation radio galaxies (LERGs), fuelled at significantly lower rates via radiatively inefficient flows. Observationally, HERGs show strong optical emission lines relative to the stellar

continuum, while LERGs show weak or no optical emission lines. In their study of the local radio source population at 20 GHz, Sadler et al. (2014) made a qualitative separation by associating Ae radio galaxies with HERGs, and Aa and Aae radio galaxies with LERGs. If we make the same separation, 10% of the AGN in the GLEAM-6dFGS sample are classified as HERGs.

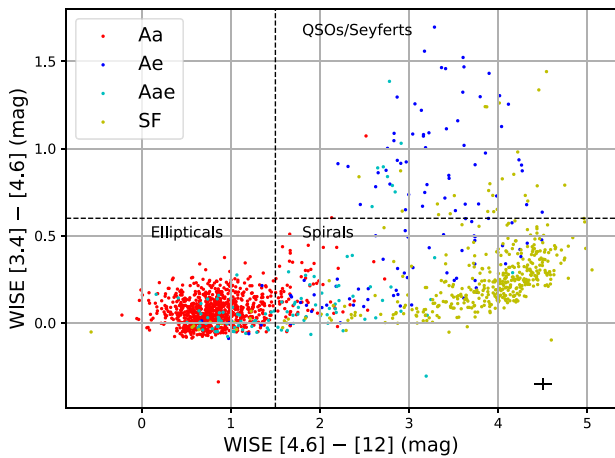
In Figure 7, we report a WISE colour-colour plot (Wright et al., 2010) in W1 (3.4  $\mu\text{m}$ ), W2 (4.6  $\mu\text{m}$ ), and W3 (12  $\mu\text{m}$ ). The median errors on W1–W2 and W2–W3 are 0.033 and 0.10, respectively. The dashed horizontal line at a W1–W2 colour of +0.6

<sup>a</sup>The reduced  $\chi^2$  value from the power-law fit can be used to assess the quality of the fitted spectral index: for 18 degrees of freedom,  $P(\text{reduced } \chi^2 > 1.93) < 1\%$  and  $P(\text{reduced } \chi^2 > 2.35) < 0.1\%$ .



**Table 5.** Spectral classes of the GLEAM-6dFGS objects.

Class	Type of spectrum	Number
Aa	Pure absorption line spectrum	933 (58.7%)
Ae	Strong narrow emission lines	118 (7.4%)
Aae	Weak narrow emission lines	106 (6.7%)
SF	HII region-like emission spectrum	427 (26.9%)
?	Unknown	6 (0.4%)
Total		1 590

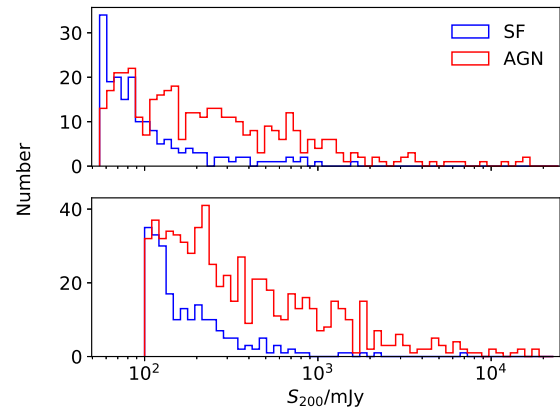
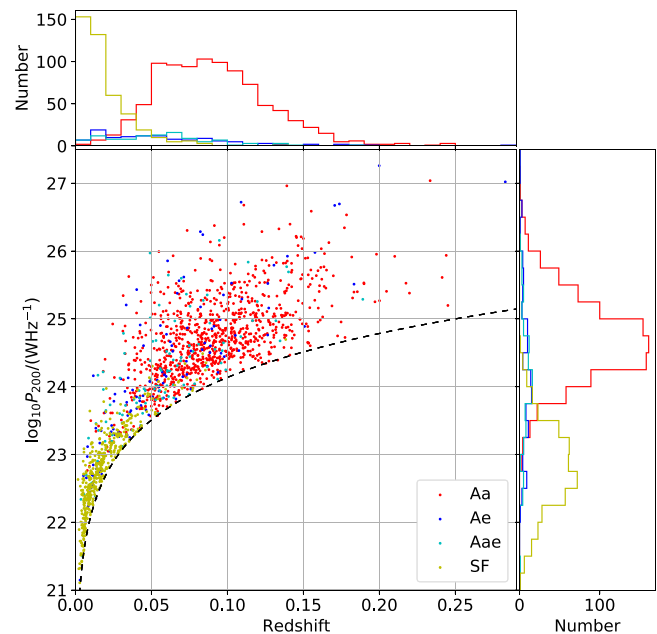
**Figure 7.** WISE colour–colour plot for the GLEAM-6dFGS galaxies, which are colour-coded based on the type of optical spectrum. The WISE magnitudes are in the Vega system. The horizontal and vertical dashed lines divide the different galaxy populations, as discussed in the text. Individual error bars are not plotted for clarity but the median error bar size for the sample is shown at the bottom right.

mag shows the dividing line between normal galaxies (LERGs) and radiatively efficient AGN (HERGs), and the vertical dashed line at a  $W2-W3$  colour of +1.5 mag shows the dividing line between elliptical and spiral galaxies, both adopted by Wright *et al.*

The WISE galaxy classes are broadly consistent with the visual classification of the optical spectra: only 2.4% of the SF galaxies lie in the ‘WISE elliptical’ region. However, 53% of the Ae galaxies lie in the ‘WISE LERG’ region. In a study of the WISE properties of a large sample of radio-identified galaxies and QSOs, Ching *et al.* (2017) noted that the Ae/HERG objects are a strongly heterogeneous class spanning a wide area in the WISE two-colour diagram. They concluded that the physical process dominating the optical/MIR light is not the same for all members of this class, and that the position of individual HERGs in the WISE two-colour plot is likely to depend on the star formation rate within the host galaxy, the brightness of the optical AGN relative to the host galaxy, and the amount of dust obscuration present within the galaxy. Our results for the 6dFGS radio-detected galaxies are consistent with this picture.

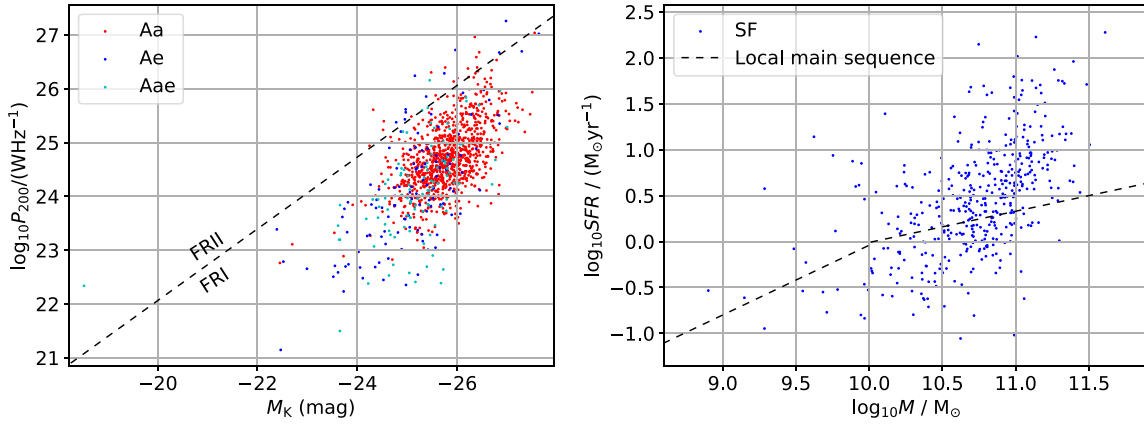
#### 4.1. Flux density distribution

Figure 8 shows the 200-MHz flux density distributions of the AGN and SF galaxies in the deep and shallow regions. The number of AGN and SF galaxies is roughly equal at 100 mJy, the flux density limit in the shallow region. The sample is dominated by SF galaxies close to the flux density limit in the deep region (55 mJy).

**Figure 8.** Histogram of  $S_{200\text{ MHz}}$  for the AGN (red) and SF galaxies (blue) in the deep (top) and shallow (bottom) regions of the GLEAM-6dFGS sample. Flux density cuts of 55 and 100 mJy are applied in the deep and shallow regions, respectively.**Figure 9.** 200-MHz radio luminosity as a function of redshift for the AGN and SF galaxies in the GLEAM-6dFGS sample. The dashed line corresponds to the 200 MHz flux density limit of 55 mJy in the deep region.

#### 4.2. Redshift and luminosity distributions

In Figure 9, we present the 200-MHz radio luminosity and redshift distributions of the AGN and SF galaxies. The AGN have a median redshift of 0.081. In comparison, the median redshift of the AGN in the NVSS-6dFGS sample by Mauch & Sadler (2007), selected at 1.4 GHz, is 0.073. The NVSS is estimated to be 99% complete at  $S_{1.4\text{ GHz}} = 3.4$  mJy. The equivalent completeness limit of NVSS at 200 MHz for a compact radio source with a spectral index of  $-0.7$  is  $\approx 13$  mJy. This is a factor of  $\approx 4-8$  lower than the 200-MHz flux density limits of the GLEAM-6dFGS sample. The fact that the median redshift of the AGN in the NVSS-6dFGS sample is similar to that in the GLEAM-6dFGS sample implies that AGN drop out of the GLEAM-6dFGS sample primarily due to the  $K$ -band magnitude limit.



**Figure 10.** Left: distribution in  $K$ -band absolute magnitude and 200-MHz radio power for the AGN in the GLEAM-6dFGS sample. The dashed line corresponds to the FRI/FRII division at 1.4 GHz from Ledlow & Owen (1996), extrapolated to  $K$ -band and 200 MHz as described in the text, with FRII galaxies lying above the line. Right: distribution in the SFR and stellar mass for the SF galaxies in the GLEAM-6dFGS sample. The stellar mass is derived from the  $K$ -band absolute magnitude and the SFR from the 200-MHz radio power and stellar mass, as described in the text. The dashed line shows the SF main sequence for galaxies in the local universe by Renzini & Peng (2015) and Popesso et al. (2019).

The SF galaxies form a more nearby population with a median redshift of 0.015. The median redshift (0.035) of the SF galaxies in the NVSS-6dFGS sample is more than twice as high as that in the GLEAM-6dFGS sample, indicating that it is mainly the radio flux density in GLEAM that limits the maximum distance to which SF galaxies can be detected.

As expected, there is a strong separation between the two source populations according to the radio luminosity, as seen in studies at 1.4 GHz. The AGN have median  $P_{200\text{ MHz}} = 10^{24.61} \text{ W Hz}^{-1}$  and the SF galaxies have median  $P_{200\text{ MHz}} = 10^{22.79} \text{ W Hz}^{-1}$ . Among the AGN, 82% have  $P_{200\text{ MHz}} > 10^{24} \text{ W Hz}^{-1}$  while 97% of the SF galaxies have  $P_{200\text{ MHz}} < 10^{24} \text{ W Hz}^{-1}$ . There is a strong increase in the fraction of sources with optical emission lines with decreasing luminosity: while only 19% of all AGN are classed as Aae or Ae, this fraction rises to 51% among the population of low-luminosity ( $P_{200\text{ MHz}} < 10^{24} \text{ W Hz}^{-1}$ ) AGN overlapping in radio luminosity with the SF galaxies.

The  $K$ -band absolute magnitude as a function of the 200-MHz radio luminosity for the AGN is shown in the left panel of Figure 10. The dashed line corresponds to the 1.4-GHz FRI/FRII dividing line from Ledlow & Owen (1996), shifted to  $K$ -band and 200 MHz assuming a typical galaxy colour of  $R-K = 3.0 \text{ mag}$  and a typical radio spectral index  $\alpha_{200\text{ MHz}}^{1400\text{ MHz}} = -0.7$ . According to this relation, our sample of nearby AGN is overwhelmingly dominated by FRI galaxies: only 29 of the 1157 AGN (2.5%) lie above the dividing line. We note, however, that there have been a number of works showing that this division line is not nearly as clean as Ledlow & Owen suggested (see e.g. Mingo et al., 2019). Secondly, our sample may contain a significant fraction of compact radio sources which are often referred to as FR0 galaxies.

Since the low-frequency radio emission from SF galaxies is primarily due to the acceleration of electrons via shocks associated with supernovae, the radio luminosity from SF galaxies can be used as a tracer of the star formation rate (SFR; e.g. Condon, 1992). Smith et al. (2021) studied the relationship between 150-MHz luminosity and SFR using LOFAR measurements at 150 MHz for a near-infrared selected sample of  $z < 1$  galaxies. We derive SFRs for the SF galaxies in our sample using their best-fit relation, which

accounts for the stellar mass,  $M_*$ :

$$\log_{10} \left( \frac{P_{150\text{ MHz}}}{\text{W Hz}^{-1}} \right) = 0.90 \log_{10} \left( \frac{\text{SFR}}{\text{M}_{\odot} \text{ yr}^{-1}} \right) + 0.33 \log_{10} \left( \frac{M_*}{10^{10} \text{ M}_{\odot}} \right) + 22.22. \quad (7)$$

We extrapolate the 200-MHz radio powers measured in this paper to 150 MHz using  $\alpha_{\text{low}}$ , if  $\Delta\alpha_{\text{low}} < 0.2$ , otherwise we use  $\alpha_{\text{high}}$ . We derive  $M_*$  from  $M_{K,\text{abs}}$  by combining equations (4) and (5) of Longhetti & Saracco (2009):

$$\log_{10} \frac{M_*}{\text{M}_{\odot}} = -0.4 (M_{K,\text{abs}} - M_{K,\text{abs,Sun}}) + \log_{10} \frac{\Upsilon_K}{\text{M}_{\odot}/L_{\odot}}, \quad (8)$$

where  $\Upsilon_K$  is the mass-to-light ratio in the  $K$ -band and  $M_{K,\text{abs,Sun}}$  is the absolute  $K$ -band magnitude of the Sun. Since disc/spiral galaxies in the local universe typically have  $\Upsilon_K \approx 0.6$  (McGaugh & Schombert, 2014) and  $M_{K,\text{abs,Sun}} \approx 3.27 \text{ mag}$  in the Vega system (Willmer, 2018):

$$\log_{10} \frac{M_*}{\text{M}_{\odot}} \approx 1.09 - 0.4 M_{K,\text{abs}}. \quad (9)$$

The SFR as a function of  $M_*$  is plotted in the right panel of Figure 10. The dashed line shows the main sequence of SF galaxies defined by Renzini & Peng (2015) at  $M_* < 10^{10} \text{ M}_{\odot}$  and by Popesso et al. (2019) at  $M_* \geq 10^{10} \text{ M}_{\odot}$ . The SF galaxies in the GLEAM-6dFGS sample scatter around and above the high-mass end of the SF main sequence. This is consistent with them lying mainly in the ‘starburst’ and ‘LIRG/ULIRG’ regions of the WISE colour–colour diagram shown in Figure 7.

### 4.3. Radio spectra

All sources in the GLEAM-6dFGS catalogue have a measurement of  $\alpha_{\text{high}}$ , while 92% of the sources have a measurement of  $\alpha_{\text{low}}$ . Figure 11 illustrates the wide variety of radio SEDs in the sample.

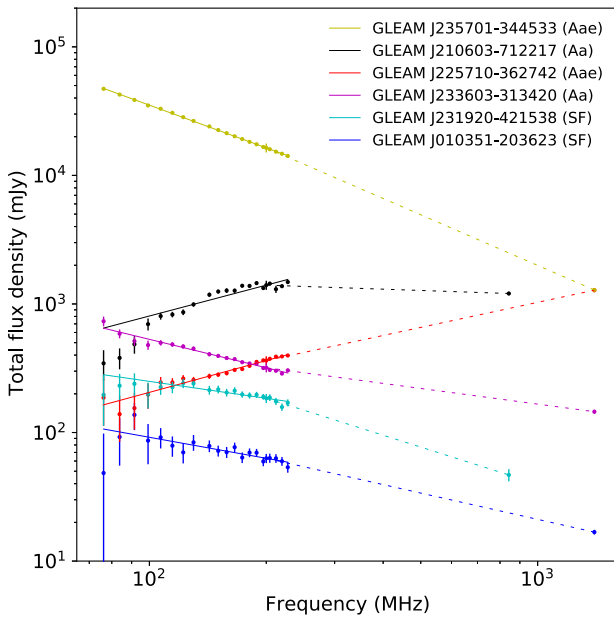
For  $\alpha_{\text{low}}$  to be quoted in the catalogue, the source must have a positive flux density in each of the 20 GLEAM sub-bands, which is not always the case at low signal-to-noise ratio (SNR). Of the 1590 sources in the GLEAM-6dFGS sample, we investigate the

**Table 6.** Spectral index statistics for the AGN and SF galaxies with measured  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  in the GLEAM-6dFGS sample.  $N_{\text{flat}}$  is the number of flat-spectrum sources with  $\alpha > -0.5$  and  $N_{\text{USS}}$  is the number of ultra-steep-spectrum sources with  $\alpha < -1.2$ .

Class	$N_{\text{total}}$	$\alpha_{\text{low}}$			$\alpha_{\text{high}}$				
		Median	Std. dev.	$N_{\text{flat}}$	$N_{\text{USS}}$	Median	Std. dev.	$N_{\text{flat}}$	$N_{\text{USS}}$
Aa	869	$-0.713 \pm 0.011$	0.26	107 (12%)	39 (4.5%)	$-0.610 \pm 0.011$	0.27	197 (23%)	44 (5.1%)
Ae	108	$-0.612 \pm 0.036$	0.30	34 (31%)	1 (0.9%)	$-0.575 \pm 0.029$	0.24	34 (31%)	2 (1.9%)
Aae	93	$-0.667 \pm 0.048$	0.37	26 (28%)	5 (5.4%)	$-0.530 \pm 0.042$	0.33	38 (41%)	3 (3.2%)
All AGN	1070	$-0.704 \pm 0.011$	0.28	167 (16%)	45 (4.2%)	$-0.600 \pm 0.010$	0.27	269 (25%)	49 (4.6%)
SF	369	$-0.596 \pm 0.015$	0.23	111 (30%)	5 (1.4%)	$-0.650 \pm 0.010$	0.15	38 (10%)	1 (0.3%)

**Table 7.** Results of a linear regression analysis to investigate the dependence of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  on  $P_{200\text{MHz}}$  for the AGN and SF galaxies in the GLEAM-6dFGS sample.  $\rho$  is the correlation coefficient. The null hypothesis that the slope is zero is evaluated using a Wald test.

Class	$\alpha_{\text{low}}$			$\alpha_{\text{high}}$		
	$\rho$	Probability of null hypothesis	$\Delta\alpha$ per log. luminosity decade	$\rho$	Probability of null hypothesis	$\Delta\alpha$ per log. luminosity decade
AGN	-0.16	$8.68 \times 10^{-8}$	$-0.062 \pm 0.011$	-0.10	$1.10 \times 10^{-3}$	$-0.037 \pm 0.011$
SF	-0.00	$9.51 \times 10^{-1}$	$-0.001 \pm 0.020$	-0.10	$6.71 \times 10^{-2}$	$-0.023 \pm 0.012$



**Figure 11.** Example SEDs using the 20 GLEAM sub-band flux densities between 76 and 227 MHz, the GLEAM wide-band flux density at 200 MHz and the NVSS/SUMSS flux density at 1400/843 MHz. For each source, the solid line shows the power-law fit to the GLEAM sub-band flux densities and the dashed line connects the GLEAM wide-band flux density to the NVSS/SUMSS flux density. Yellow: AGN classed as Aae with a steep power-law spectrum across the GLEAM band; black: AGN classed as Aa peaking at around 200 MHz; red: AGN classed as Aae with a highly inverted spectrum; magenta: AGN classed as Aa with a flattening spectrum at high frequency; cyan: SF galaxy with a turnover frequency around 100 MHz; blue: SF galaxy close to the 200-MHz detection limit of the sample with a typical SED.

spectral properties of the sub-sample of 1 439 sources that have measured  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ , excluding the 24 sources severely affected by confusion in GLEAM and the six sources with ‘unknown’

optical spectra. We note that this sub-sample is biased against faint sources with flat spectra across the GLEAM frequency range: the negative flux densities mostly occur in the lowest frequency sub-bands which have higher rms noise. It is also biased against diffuse radio sources with ultra-steep radio spectra, as discussed in Section 3.2. The median errors on  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  are 0.08 and 0.06, respectively.

The left panel of Figure 12 shows a radio colour–colour plot for all AGN in the sub-sample, which are colour-coded according to their optical spectral classification. The right panel shows a radio colour–colour plot for all SF galaxies in the sub-sample. In Table 6, the median and standard deviation values of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  are provided for each class of optical spectrum. Table 6 also gives the fractions of sources with  $\alpha_{\text{low}} > -0.5$ ,  $\alpha_{\text{high}} > -0.5$ ,  $\alpha_{\text{low}} < -1.2$  and  $\alpha_{\text{high}} < -1.2$ .

#### 4.3.1. AGN

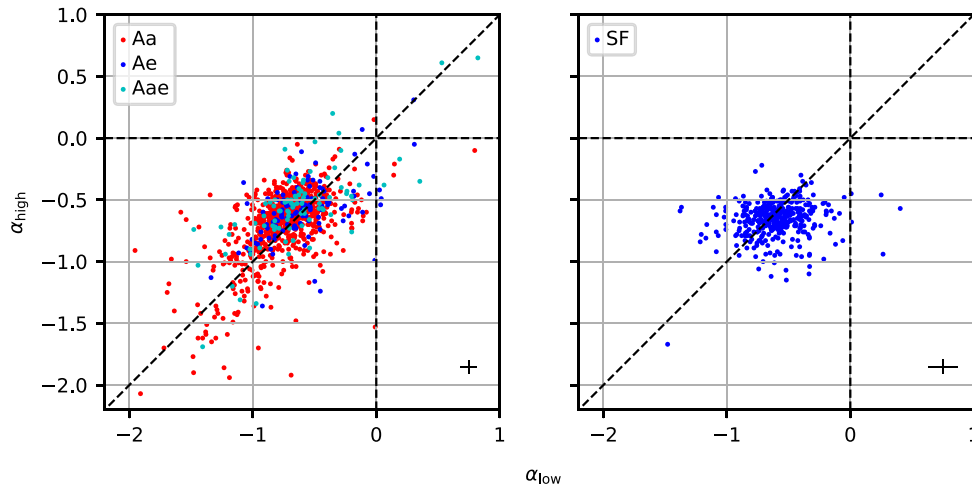
For the AGN, the median value of  $\alpha_{\text{low}}$  is  $-0.704 \pm 0.011$  and the median value of  $\alpha_{\text{high}}$  is  $-0.600 \pm 0.010$ . The slight spectral flattening with frequency is likely due to the radio emission becoming increasingly core-dominated at high frequency. In their LoTSS-SDSS sample of nearby galaxies, Sabater *et al.* (2019) measured the 150–1 400 MHz spectral indices of 496 radio AGN detected in NVSS and FIRST. For sources with  $S_{1400\text{MHz}} > 20$  mJy, they obtained a median spectral index of  $-0.63$ , which is close to the median value of  $\alpha_{\text{high}}$  for the AGN in the GLEAM-6dFGS sample. The comparison is performed over a similar flux density range since  $\approx 90\%$  of the AGN in the GLEAM-6dFGS sample have NVSS/SUMSS flux densities above 20 mJy.

For the G4Jy sample, which extends to much higher redshifts and luminosities than the GLEAM-6dFGS sample, the spectral indices over the same frequency ranges are steeper by  $\sim 0.15$ : the median value of  $\alpha$  is  $-0.740 \pm 0.012$  between 151 and 843 MHz,  $-0.786 \pm 0.006$  between 151 and 1 400 MHz, and  $-0.829 \pm 0.006$



**Table 8.** Local RLFs at 200 MHz for all radio sources, the AGN and SF galaxies in the GLEAM-6dFGS sample.

$\log_{10} P_{200 \text{ MHz}}$ (W Hz <sup>-1</sup> )	All galaxies		AGN		SF galaxies	
	$N$	$\log_{10} \Phi$ (mag <sup>-1</sup> Mpc <sup>-3</sup> )	$N$	$\log_{10} \Phi$ (mag <sup>-1</sup> Mpc <sup>-3</sup> )	$N$	$\log_{10} \Phi$ (mag <sup>-1</sup> Mpc <sup>-3</sup> )
21.8	32	-2.94 <sup>+0.07</sup> <sub>-0.09</sub>	0		32	-2.94 <sup>+0.07</sup> <sub>-0.09</sub>
22.2	65	-3.13 <sup>+0.07</sup> <sub>-0.09</sub>	5	-3.88 <sup>+0.25</sup> <sub>-0.70</sub>	60	-3.21 <sup>+0.06</sup> <sub>-0.06</sub>
22.6	131	-3.42 <sup>+0.04</sup> <sub>-0.04</sub>	20	-4.29 <sup>+0.09</sup> <sub>-0.12</sub>	111	-3.49 <sup>+0.04</sup> <sub>-0.05</sub>
23.0	121	-4.04 <sup>+0.04</sup> <sub>-0.05</sub>	17	-4.95 <sup>+0.10</sup> <sub>-0.13</sub>	104	-4.09 <sup>+0.04</sup> <sub>-0.05</sub>
23.4	106	-4.64 <sup>+0.04</sup> <sub>-0.05</sub>	41	-5.06 <sup>+0.07</sup> <sub>-0.08</sub>	65	-4.86 <sup>+0.05</sup> <sub>-0.06</sub>
23.8	147	-5.11 <sup>+0.04</sup> <sub>-0.04</sub>	118	-5.21 <sup>+0.04</sup> <sub>-0.05</sub>	29	-5.76 <sup>+0.08</sup> <sub>-0.09</sub>
24.2	233	-5.38 <sup>+0.03</sup> <sub>-0.03</sub>	221	-5.41 <sup>+0.03</sup> <sub>-0.03</sub>	12	-6.54 <sup>+0.12</sup> <sub>-0.17</sub>
24.6	306	-5.54 <sup>+0.03</sup> <sub>-0.03</sub>	304	-5.54 <sup>+0.03</sup> <sub>-0.03</sub>	2	-7.72 <sup>+0.36</sup> <sub>-0.45</sub>
25.0	211	-5.81 <sup>+0.04</sup> <sub>-0.04</sub>	211	-5.81 <sup>+0.04</sup> <sub>-0.04</sub>	0	
25.4	132	-6.07 <sup>+0.05</sup> <sub>-0.06</sub>	132	-6.07 <sup>+0.05</sup> <sub>-0.06</sub>	0	
26.8	60	-6.53 <sup>+0.09</sup> <sub>-0.11</sub>	60	-6.53 <sup>+0.09</sup> <sub>-0.11</sub>	0	
26.2	19	-7.06 <sup>+0.11</sup> <sub>-0.15</sub>	19	-7.06 <sup>+0.11</sup> <sub>-0.15</sub>	0	
26.6	8	-7.62 <sup>+0.15</sup> <sub>-0.23</sub>	8	-7.62 <sup>+0.15</sup> <sub>-0.23</sub>	0	
27.2	4	-8.59 <sup>+0.25</sup> <sub>-0.28</sub>	4	-8.59 <sup>+0.25</sup> <sub>-0.28</sub>	0	
Total	1575		1160		415	



**Figure 12.** Left: radio colour-colour diagram for all AGN in the GLEAM-6dFGS sample with measured  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ ;  $\alpha_{\text{low}}$  is the spectral index between 76 and 227 MHz, and  $\alpha_{\text{high}}$  is the spectral index between 200 and 1400/843 MHz. AGN classed as Aa are shown in red, Ae in blue and Aae in cyan. The dashed lines represent spectral indices of zero and equal values of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ . Individual error bars are not plotted for clarity, but the median error bar size for the sample is shown at the bottom right. Right: radio colour-colour diagram for all SF galaxies in the GLEAM-6dFGS sample with measured  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ .

within the GLEAM band (White et al. 2020a). In Figure 13, we investigate the dependence of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  on  $P_{200 \text{ MHz}}$  within the GLEAM-6dFGS sample, separating between the AGN and SF galaxies. The results of a linear regression analysis to quantify any correlation between the spectral index and radio luminosity for both populations are given in Table 7. For the AGN, there is a statistically significant anti-correlation between  $\alpha_{\text{low}}$  and  $P_{200 \text{ MHz}}$ , and  $\alpha_{\text{high}}$  and  $P_{200 \text{ MHz}}$ ;  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  steepen by  $0.062 \pm 0.011$  and  $0.037 \pm 0.011$  per logarithmic luminosity decade, respectively. For the SF galaxies, there is no statistically significant correlation between  $\alpha_{\text{low}}$  and  $P_{200 \text{ MHz}}$ , and  $\alpha_{\text{high}}$  and  $P_{200 \text{ MHz}}$ .

Laing & Peacock (1980) investigated the relation between spectral index and radio luminosity for a complete sample of 165

sources with  $S_{178 \text{ MHz}} > 10 \text{ Jy}$ , consisting primarily of FR II sources. For the extended structure in the FR II sources,  $\alpha_{0.75 \text{ GHz}}^{5 \text{ GHz}}$  was found to be anti-correlated with the 1.4-GHz radio luminosity. The slope of the regression line (i.e. the change in the spectral index per logarithmic luminosity decade) was measured as  $-0.088$ . A similar anti-correlation was found between  $\alpha_{0.75 \text{ GHz}}^{5 \text{ GHz}}$  and redshift, which is expected given the strong correlation between luminosity and redshift in bright, flux-limited samples. The  $P - \alpha$  anti-correlation was confirmed by Blundell, Rawlings, & Willott (1999) for a number of complete samples of FR II sources selected at frequencies close to 151 MHz. They suggested that the anti-correlation is caused by radiative losses in the enhanced magnetic fields of the hotspots of sources with more powerful jets.

We find that the  $P - \alpha$  anti-correlation continues to hold for less powerful sources in the local universe. de Gasperin, Intema, & Frail (2018) investigated the spectral index properties between 147 and 1 400 MHz of radio sources from TGSS and NVSS, and reported a flattening of the median spectral index with decreasing flux density. A similar spectral flattening towards lower flux densities has been observed in deeper radio samples covering smaller sky areas (e.g. Prandoni *et al.*, 2006; Intema *et al.*, 2011; Whittam *et al.*, 2013). Since compact radio cores display much flatter spectra and are typically much fainter than the lobes of well-evolved radio galaxies, de Gasperin *et al.* argued that this trend is driven by the presence of core-dominated and young AGN at low flux densities. Our results for the GLEAM-6dFGS AGN are consistent with this picture. Flat-spectrum sources are more common at lower radio luminosity, suggesting the existence of a significant population of low-luminosity AGN that remain core-dominated even at low frequencies.

Figure 14 shows that there is a strong increase in the fraction of sources with optical emission lines with  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ . In other words, optical emission lines are much more common in galaxies which host flat-spectrum radio sources. In their study of the local radio source population at 20 GHz, Sadler *et al.* (2014) observed the same trend for radio galaxies with weak optical emission lines (class Aae) when measuring spectral indices between  $\sim 1$  and 20 GHz. The fraction of radio galaxies with strong emission lines (class Ae) was not found to change significantly with spectral index.

Zajaček *et al.* (2019) studied radio spectral index trends of SDSS-FIRST sources across optical emission line diagnostic diagrams to search for potential correlations. They found a trend of decreasing 4.85–10.45 GHz spectral index with increasing ionisation ratio [OIII]/H $\beta$ , which corresponds to the LINER-Seyfert transition in the optical diagnostic diagrams. They interpreted this trend as a result of re-triggered nuclear and jet activity.

In low-frequency selected samples, there is preferential detection of sources with ultra-steep spectra down to low frequencies. A significant fraction of AGN in the GLEAM-6dFGS sample have ultra-steep spectra: 4.2% have  $\alpha_{\text{low}} < -1.2$  and 4.6% have  $\alpha_{\text{high}} < -1.2$ . Mahatma *et al.* (2018) reported a similar fraction (4.1%) of sources with  $\alpha_{150}^{1400} < -1.2$  in their LOFAR study of remnant radio-loud AGN in the Herschel-ATLAS field. The ultra-steep-spectrum AGN in our sample have a wide range of radio luminosities; the median value of  $P_{200 \text{ MHz}}$  for the AGN with  $\alpha_{\text{low}} < -1.2$  ( $10^{24.41} \text{ W Hz}^{-1}$ ) is slightly lower than the median value of  $P_{200 \text{ MHz}}$  for all AGN ( $10^{24.61} \text{ W Hz}^{-1}$ ). Quici *et al.* (2021) show that remnant radio galaxies can be reliably identified by their ultra-steep spectra at low frequency ( $\lesssim 400 \text{ MHz}$ ), although the technique fails to capture younger remnants with steeper spectra at high frequencies. The absence of a radio core at high frequency is needed to establish whether these are genuine remnant radio galaxies.

### 4.3.2. SF galaxies

For the SF galaxies, there is a slight flattening in the typical spectral index towards lower frequency: the median value of  $\alpha_{\text{low}} = -0.596 \pm 0.015$  and the median value of  $\alpha_{\text{high}} = -0.650 \pm 0.010$ . A larger fraction of SF galaxies have flat spectra at low frequency: 10% have  $\alpha_{\text{high}} > -0.5$  while 30% have  $\alpha_{\text{low}} > -0.5$ . We also find

that the distribution of  $\alpha_{\text{low}}$  is significantly broader than that of  $\alpha_{\text{high}}$ . We are able to show that this result is robust and not just a consequence of the larger measurement errors on  $\alpha_{\text{low}}$ . Using Monte Carlo simulations which account for the measurement errors, the intrinsic dispersions in  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  are estimated to be  $0.186 \pm 0.012$  and  $0.127 \pm 0.008$ , respectively. These simulations are described in detail in Appendix A.

Chyży *et al.* (2018) also found a flattening in the typical spectral index towards lower frequency for a sample of 106 nearby SF galaxies using measurements over a larger frequency range. Using 150-MHz flux density measurements from the LOFAR Multifrequency Snapshot Sky Survey (MSSS; Heald *et al.*, 2015) and literature flux densities at various frequencies, they obtained a median value of  $\alpha_{50 \text{ MHz}}^{1.5 \text{ GHz}} = -0.57$  and a median value of  $\alpha_{1.3 \text{ GHz}}^5 = -0.77$ . Since there was no tendency for the highly inclined galaxies to have flatter low-frequency spectra, they argued that the observed flattening was not due to thermal absorption. From numerical modelling of the radio emission, they inferred that the flattening resulted principally from synchrotron spectral curvature due to cosmic ray energy losses and propagation effects.

## 5. The local RLF at 200 MHz

### 5.1. Calculating the local RLF

We calculate the local RLF using the  $1/V_{\text{max}}$  method of Schmidt (1968), where  $V_{\text{max}}$  is the maximum volume within which a galaxy can satisfy all the sample selection criteria ( $S_{200 \text{ MHz}} > 100 \text{ mJy}$  in the shallow region,  $S_{200 \text{ MHz}} > 55 \text{ mJy}$  in the deep region, and  $K \leq 12.65 \text{ mag}$  in both the shallow and deep regions). The luminosity function in a given luminosity bin centred at  $L$  is given by:

$$\Phi(L) = \sum_{i=1}^N \left( \frac{1}{V_{\text{max, deep}, i} + V_{\text{max, shallow}, i}} \right), \quad (10)$$

where  $N$  is the number of galaxies in the luminosity bin, and  $V_{\text{max, deep}, i}$  and  $V_{\text{max, shallow}, i}$  are the maximum volumes within which the  $i^{\text{th}}$  galaxy can satisfy the sample selection criteria in the deep and shallow regions respectively. The rms Poisson counting error in  $\Phi(L)$  is given by:

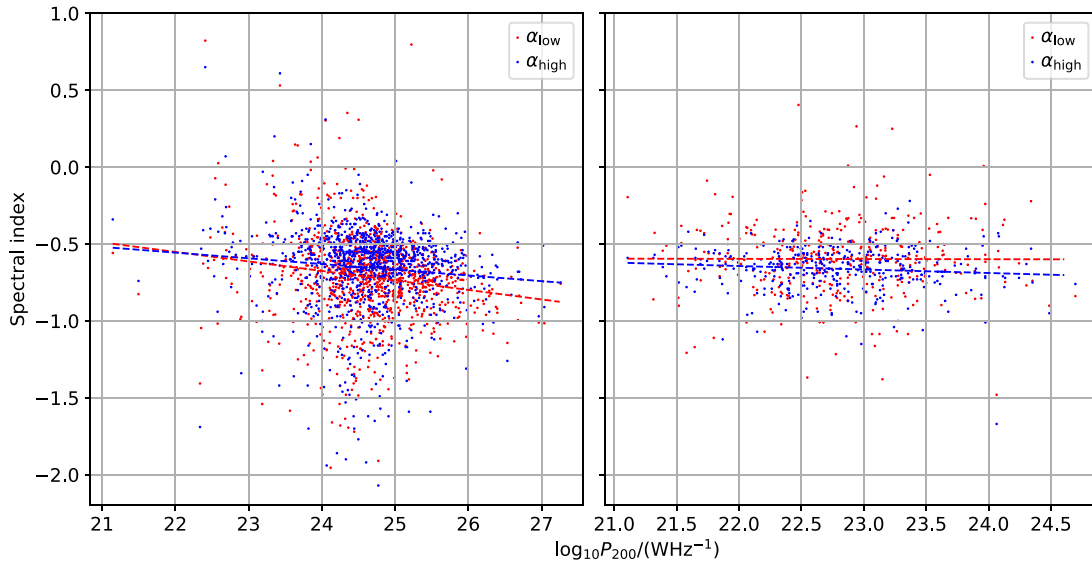
$$d\Phi(L) = \left[ \sum_{i=1}^N \left( \frac{1}{V_{\text{max, deep}, i} + V_{\text{max, shallow}, i}} \right)^2 \right]^{1/2}. \quad (11)$$

We calculate the luminosity function in 13 luminosity bins of width 1 mag (= 0.4 dex) between  $10^{21.6}$  and  $10^{26.8} \text{ W Hz}^{-1}$ , and one luminosity bin of width 2 mag between  $10^{26.8}$  and  $10^{27.6} \text{ W Hz}^{-1}$ . In luminosity bins where the number,  $N$ , of galaxies is small ( $N < 5$ ), we use 84% confidence upper and lower limits based on Poisson statistics, tabulated in Gehrels (1986).

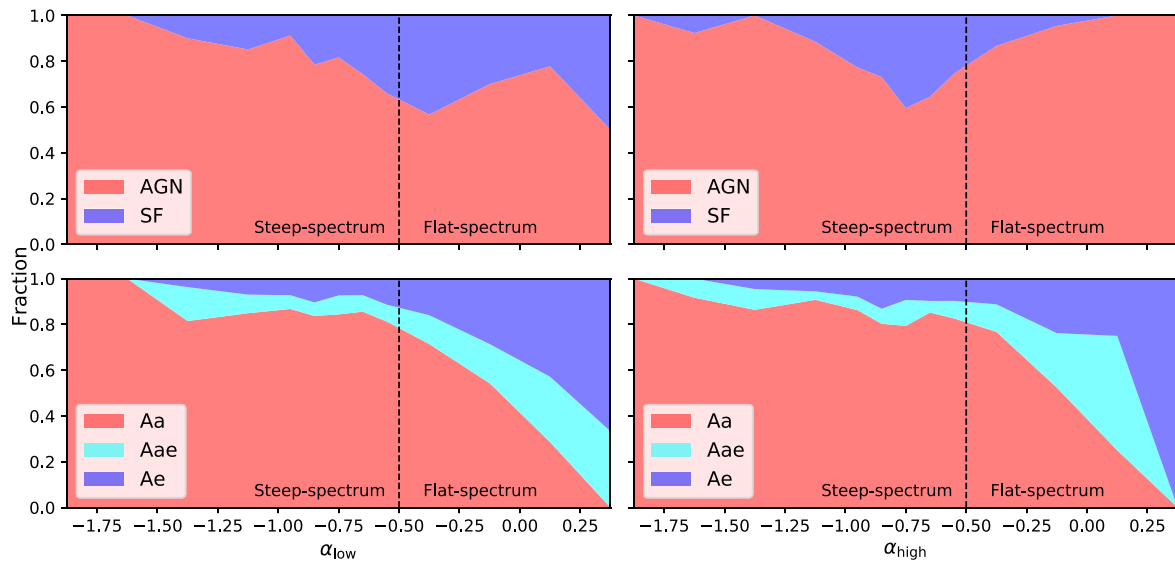
Following Condon *et al.* (2019), we correct the measured luminosity function for galaxy clustering at a distance  $s$  centred on our own Galaxy using

$$\frac{\Phi_p}{\Phi} = 1 + \frac{3}{3 - \gamma_s} \left( \frac{s_0}{s} \right)^{\gamma_s}, \quad (12)$$

where  $\frac{\Phi_p}{\Phi}$  is the expected overdensity near our own Galaxy, or the space density,  $\Phi_p$ , of local galaxies divided by the average space density,  $\Phi$ , of all galaxies,  $\gamma_s = 1.66$  and  $s_0 = 3.76h^{-1} \text{ Mpc}$  (Fisher *et al.*, 1994). The volume within  $s$  is multiplied



**Figure 13.**  $\alpha_{low}$  (red) and  $\alpha_{high}$  (blue) as a function of the 200-MHz radio luminosity for the AGN (left) and SF galaxies (right) in the GLEAM-6dFGS sample with measured  $\alpha_{low}$  and  $\alpha_{high}$ . The dashed lines are the linear regression lines.



**Figure 14.** Top: fractions of AGN and SF galaxies as a function of  $\alpha_{low}$  (left) and  $\alpha_{high}$  (right). Bottom: among the AGN, fractions of Aa, Aae, and Ae sources as a function of  $\alpha_{low}$  (left) and  $\alpha_{high}$  (right).

by  $\frac{\Phi_p}{\Phi}$  when calculating  $V_{max}$ ; the correction is only applied for  $s < 20h^{-1}$  Mpc.

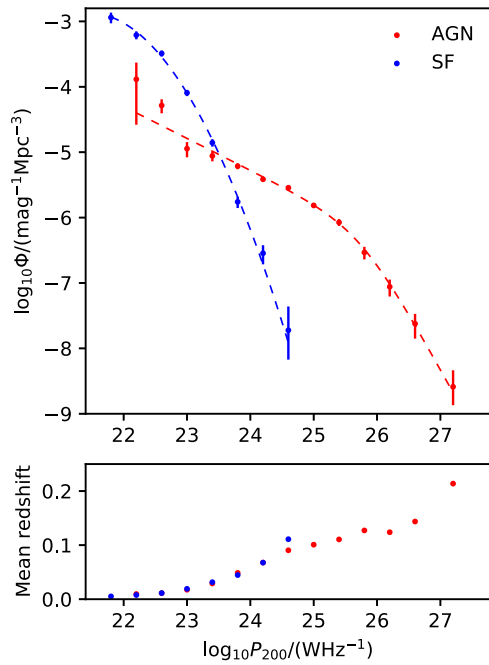
The completeness of the GLEAM-6dFGS cross-matched sample is estimated to be 95% (see Section 3.1.1). To correct the luminosity function for incompleteness, we therefore multiply  $\Phi(L)$  in all bins by  $\frac{1}{0.95}$ . Additionally, we need to account for spectroscopic incompleteness in the 6dFGS catalogue, that is, the fact that some galaxies in the 6dFGS target list were either not able to be observed or had only a poor-quality spectrum so that a reliable redshift could not be measured. The 6dFGS spectroscopic completeness varies across the sky but is estimated to be 85% for the NVSS/SUMSS-6dFGS sample as a whole. To

correct the luminosity function for spectroscopic incompleteness, we therefore multiply  $\Phi(L)$  in all bins by a further factor of  $\frac{1}{0.85}$ .

**5.2. Results**

Our measured local RLF at 200 MHz for AGN and SF galaxies is tabulated in Table 8 and plotted in the top panel of Figure 15. The six sources with unknown optical spectra are classified according to their 200-MHz luminosity: those with  $P_{200\text{ MHz}} > 10^{23.5} \text{ W Hz}^{-1}$  are classed as AGN and those with  $P_{200\text{ MHz}} \leq 10^{23.5} \text{ W Hz}^{-1}$  are classed as SF galaxies. In total, 3 AGN and 44 SF galaxies are





**Figure 15.** Top: the local RLFs at 200 MHz for the AGN (red) and SF galaxies (blue) in the GLEAM-6dFGS sample. The two dashed curves correspond to the double power-law and Saunders et al. (1990) fits to the AGN and SF data, respectively. Bottom: the mean redshifts of the AGN and SF galaxies in each bin.

affected by the local volume correction described above. This correction only significantly affects the measured luminosity function of the SF galaxies in the lowest luminosity bin.

The RLF for the GLEAM-6dFGS AGN is well fitted using a double power-law function of the form:

$$\Phi(L) = \frac{C}{(L/L_*)^\alpha + (L/L_*)^\beta}. \quad (13)$$

The best-fitting parameters of equation (13) are  $C = 10^{-6.13} \text{ mag}^{-1} \text{ Mpc}^{-3}$ ,  $L_* = 10^{25.76} \text{ W Hz}^{-1}$ ,  $\alpha = 1.76$ , and  $\beta = 0.49$ , where  $\alpha$  is the power-law slope for  $L \gg L_*$  and  $\beta$  is the power-law slope for  $L \ll L_*$ .

The RLF for SF galaxies is commonly fitted by the parametric form given by:

$$\Phi(L) = C \left( \frac{L}{L_*} \right)^{1-\alpha} \exp \left\{ -\frac{1}{2} \left[ \frac{\log_{10}(1 + L/L_*)}{\sigma} \right]^2 \right\}, \quad (14)$$

analogous to the far-infrared luminosity function of *Infrared Astronomical Satellite (IRAS)* galaxies (Saunders et al., 1990). This function approaches a power law at  $L \ll L_*$  and falls like a Gaussian at  $L \gg L_*$ . The high-luminosity end of the function is well sampled by the GLEAM-6dFGS SF galaxies. The best-fitting parameters of equation (14) are  $C = 10^{-2.84} \text{ mag}^{-1} \text{ Mpc}^{-3}$ ,  $L_* = 10^{21.06} \text{ W Hz}^{-1}$ ,  $\alpha = 0.68$ , and  $\sigma = 0.66$ .

While the AGN/SF galaxy classification from the optical spectra is generally reliable, there are a couple of cases in which an AGN can be misclassified as a SF galaxy or vice versa: firstly, the galaxy has a radio-quiet optical AGN, but the radio emission comes mainly from star formation (e.g. Kauffmann et al., 2003). In nearby galaxies, the main SF regions may also fall outside the region covered by the optical fibre used for spectroscopy, aggravating this problem. Secondly, the optical spectrum is SF-like, but

the radio emission comes mainly from a central AGN. Here, the AGN may be dust-obscured in the optical, as may be the case in ULIRGs (e.g. Norris et al., 2013), or else the radio emission could be a genuine composite of AGN and SF emission. Since SF galaxies outnumber AGN at  $P_{200 \text{ MHz}} \lesssim 10^{23.5} \text{ W Hz}^{-1}$ , it is more likely that a genuine SF object is misclassified as an AGN rather than the other way around. This could potentially account for part of the AGN upturn seen in the two lowest-luminosity bins in Figure 15.

The bottom panel of Figure 15 shows the mean redshifts of the AGN and SF galaxies in each luminosity bin. As expected, the mean redshift increases with luminosity due to the Malmquist bias. We note that the small change in the redshift distribution with luminosity will affect the shape of the luminosity function due to redshift evolution of the AGN and SF populations.

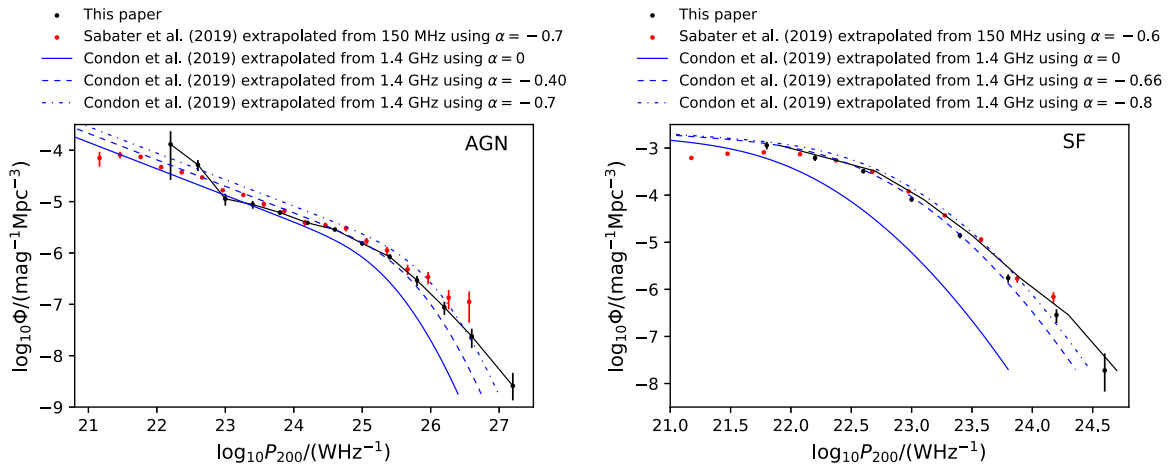
### 5.3. Comparison of recent RLF measurements

In the left panel of Figure 16, we compare the local RLF for AGN measured in this paper with the local RLF at 150 MHz for AGN from the LoTSS-SDSS sample (Sabater et al., 2019), extrapolated to 200 MHz assuming  $\alpha = -0.7$ , the typical spectral index at low frequency of AGN measured in this paper. Overall, the luminosity functions are in good agreement. The median redshift (0.143) of the AGN in the LoTSS-SDSS sample is higher than that (0.081) in the GLEAM-6dFGS sample. Cosmological evolution of HERGs, which dominate the local AGN population at  $P_{1.4 \text{ GHz}} \gtrsim 10^{26} \text{ W Hz}^{-1}$  (see e.g. Heckman & Best, 2014; Pracy et al., 2016), may therefore partly explain the slight offset between the luminosity functions at the high-luminosity end.

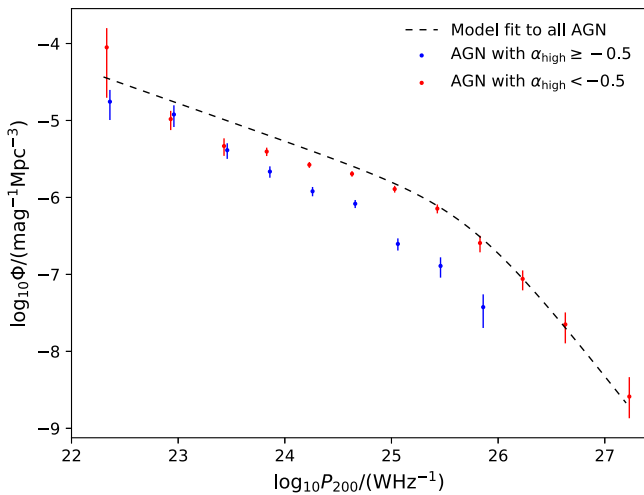
We use the parameterised form of the 1 400-MHz local RLF for AGN from equation (28) of Condon et al. (2019) and fit this to the 200-MHz data by making a shift in radio power set by a single characteristic 200–1 400 MHz spectral index. The best-fitting value of  $\alpha$  is  $-0.40$ , which is flatter than the median value of  $\alpha_{\text{high}}$  ( $-0.60$ ) for the AGN in the GLEAM-6dFGS sample. The AGN sample by Condon et al. (2019) has a slightly higher median redshift (0.12). Any luminosity evolution of the AGN population would therefore cause the best-fitting value of  $\alpha$  to be flatter than expected. We note, however, that the local AGN population is dominated by LERGs at  $P_{1.4 \text{ GHz}} \lesssim 10^{26} \text{ W Hz}^{-1}$ , which display little or no evolution (see e.g. Heckman & Best, 2014; Pracy et al., 2016). There is a significant change in the shape of the luminosity function. At lower luminosity, a flatter spectral index is needed to obtain a good fit. This is consistent with our earlier investigation of spectral index versus luminosity in Section 4.3, where we found that the typical value of  $\alpha_{\text{high}}$  flattens with decreasing luminosity.

In Figure 17, we measure the local RLF for the AGN in the GLEAM-6dFGS sample, separating between flat-spectrum ( $\alpha_{\text{high}} > -0.5$ ) and steep-spectrum ( $\alpha_{\text{high}} < -0.5$ ) sources. While steep-spectrum sources overwhelmingly dominate the luminosity function at the highest luminosities, the relative contribution to the luminosity function from flat-spectrum sources gradually increases with decreasing luminosity.

In the right panel of Figure 16, we compare the local RLF for SF galaxies measured in this paper with the local RLF at 150 MHz for SF galaxies by Sabater et al. (2019), extrapolated to 200 MHz assuming  $\alpha = -0.6$ , the typical spectral index at low frequency of SF galaxies measured in this paper. The luminosity function by Sabater et al. is shifted by  $\log_{10} P_{200 \text{ MHz}} \approx 0.1$  to the right relative to our measurements. As mentioned in Section 1, the LoTSS-SDSS sample covers a much smaller area of sky ( $424 \text{ deg}^2$ ) and extends



**Figure 16.** Left: the local RLF at 200 MHz for AGN measured in this paper (black circles). The local RLF at 150 MHz for AGN from Sabater et al. (2019) extrapolated to 200 MHz assuming a spectral index of  $-0.7$  (red circles). The local RLF at 1400 MHz for AGN using the parameterisation of Condon et al. (2019), extrapolated to 200 MHz assuming spectral indices of 0 (solid blue line),  $-0.40$  (dashed blue line), and  $-0.7$  (dot-dashed blue line), as discussed in the text. Right: the local RLF at 200 MHz for SF galaxies measured in this paper (black circles). The local RLF at 150 MHz for SF galaxies from Sabater et al. extrapolated to 200 MHz assuming a spectral index of  $-0.6$  (red circles). The local RLF at 1400 MHz for SF galaxies using the parameterisation of Condon et al. (2019), extrapolated to 200 MHz assuming spectral indices of 0 (solid blue line),  $-0.66$  (dashed blue line), and  $-0.8$  (dot-dashed blue line), as discussed in the text.



**Figure 17.** The local RLF for the GLEAM-6dFGS AGN, separating between flat-spectrum (blue points) and steep-spectrum (red points) sources. The dashed curve shows the double power-law fit to the local RLF for all AGN. There is an increase in the relative contribution of flat-spectrum sources with decreasing luminosity.

to much lower flux densities ( $\approx 0.4$  mJy) than the GLEAM-6dFGS sample. Consequently, the median redshift (0.097) of the SF galaxies in the LoTSS-SDSS sample is significantly higher than that (0.015) in the GLEAM-6dFGS sample. Ocran et al. (2020) studied the properties of SF galaxies selected at 610 MHz with the GMRT in a deep survey covering  $\approx 1.86$  deg<sup>2</sup>. They explored the evolution of the RLF for SF galaxies out to  $z \approx 1.5$ . The SF galaxies were found to evolve as  $f(z) = (1+z)^{(2.95 \pm 0.19) - (0.50 \pm 0.15)z}$  assuming pure luminosity evolution. The luminosity function by Sabater et al. is therefore expected to be shifted to the right by  $\log_{10} P_{200\text{MHz}} \approx \log_{10} [f(0.097)/f(0.015)] = 0.10$  relative to our measurements. We conclude that the luminosity functions are in good agreement after taking the rapid evolution of the SF galaxy population into account.

We use the parameterised form of the 1400-MHz local RLF for SF galaxies from equation (28) of Condon et al. (2019) and fit this to the 200-MHz data using a single characteristic 200–1400 MHz spectral index. The median redshift of the SF galaxies in the NVSS-2MASX sample by Condon et al. is 0.06. The best-fitting value of  $\alpha$  is  $-0.66$ , which is very close to the median value of  $\alpha_{\text{high}}$  ( $-0.65$ ) for the SF galaxies in our sample.

### 6. Summary and future work

We have studied the local radio source population at 200 MHz over most of the southern sky (16 679 deg<sup>2</sup>) by combining the GLEAM Exgal and SGP data with optical spectroscopy from the 6dFGS DR3. The GLEAM-6dFGS sample contains 1 590 sources with a median redshift of 0.064 and 200 MHz radio luminosities in the range  $\sim 10^{22} - 10^{27}$  W Hz<sup>-1</sup>. The optical spectra indicate that 73% of the sources are fuelled by AGN activity and 27% by star formation. The vast majority (81%) of the AGN have pure absorption-line spectra typical of giant elliptical galaxies. The AGN with emission-line spectra typically have lower luminosities and overlap in luminosity with the SF galaxies.

We characterise the typical spectra of the AGN and SF galaxies using the intra-band GLEAM spectral indices between 76 and 227 MHz, and spectral indices between 200-MHz and  $\sim 1$  GHz measured by cross-matching the sample with NVSS and SUMSS. The relatively low spatial resolution of GLEAM, NVSS, and SUMSS results in a high surface brightness sensitivity that is required to recover the extended emission in nearby galaxies and obtain accurate spectral indices.

For the AGN, the median value of  $\alpha_{\text{high}}$  ( $-0.600 \pm 0.010$ ) is significantly flatter than the median value of  $\alpha_{\text{low}}$  ( $-0.704 \pm 0.011$ ). This is likely the result of the cores of FRI sources becoming increasingly dominant at high frequencies. The 200-MHz luminosity is anti-correlated with  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ ; both  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  steepen by  $\approx 0.05$  per logarithmic luminosity decade. A significant fraction ( $\approx 4\%$ ) of the AGN in our low-frequency selected sample have ultra-steep spectra at low frequency with  $\alpha_{\text{low}} < -1.2$ . These sources span a wide range of radio luminosities. Further work is

needed to establish their nature. Optical emission lines are much more common in galaxies which host flat-spectrum radio sources, confirming the trend observed by Sadler *et al.* (2014) at higher frequency for radio galaxies in the AT20G-6dFGS sample.

For the SF galaxies, the median value of  $\alpha_{\text{low}}$  is  $\alpha = -0.596 \pm 0.015$  and the median value of  $\alpha_{\text{high}}$  is  $\alpha = -0.650 \pm 0.010$ . A larger fraction of SF galaxies have flat spectra at low frequency: 10% have  $\alpha_{\text{high}} > -0.5$  while 30% have  $\alpha_{\text{low}} > -0.5$ . Unlike for the AGN population, there is no statistically significant correlation between the 200-MHz luminosity and  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$ . Further work is needed to identify the physical processes responsible for the spectral flattening at lower frequencies.

We derive the local RLF for AGN and SF galaxies at 200 MHz. A spectral index of  $-0.40$  provides the best match between our local RLF for AGN and that measured by Condon *et al.* (2019) at 1.4 GHz. This is significantly flatter than the median value of  $\alpha_{\text{high}}$  for the AGN in the GLEAM-6dFGS sample. The difference is likely due to the changing nature of the sources contributing to the local RLF at 200 and 1 400 MHz, and the preferential selection of sources with flatter spectra at 1 400 MHz. There appears to be a slight change in the shape of the local RLF for AGN with frequency due to an increase in the relative contribution of flat-spectrum sources ( $\alpha_{\text{high}} > -0.5$ ) at lower luminosity, although this effect is difficult to measure in the lowest one or two luminosity bins due to a potential misclassification of SF galaxies as AGN. A spectral index of  $-0.66$  provides the best match between our local RLF for SF galaxies and that measured by Condon *et al.* at 1.4 GHz, which is very close to the median value of  $\alpha_{\text{high}}$  for the SF galaxies in the GLEAM-6dFGS sample.

Seymour *et al.*, in preparation, are measuring the RLF at 200 MHz out to  $z \approx 0.5$  by combining MWA data with deep multi-wavelength surveys in CDFS and ELAIS-S1. The local RLF presented in this paper will provide an essential benchmark from which to analyse the cosmic evolution of radio galaxies to this redshift.

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**Conflicts of interest.** None.

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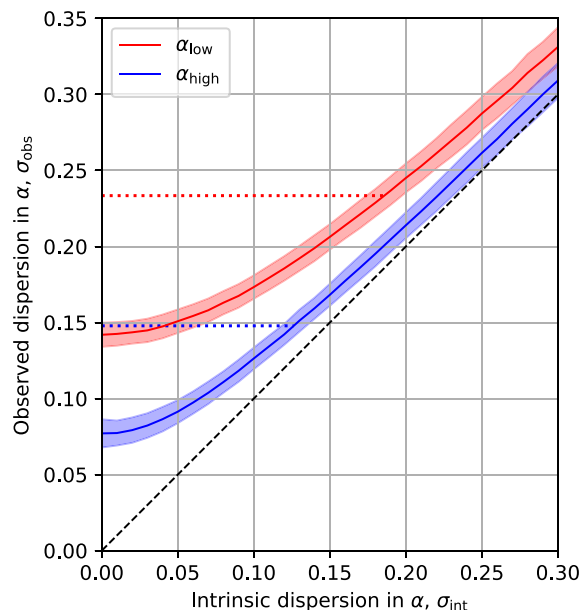
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### A. Simulations to estimate the intrinsic spectral index dispersion for the SF galaxies

As part of the spectral index analysis in Section 4.3, we perform Monte Carlo simulations to estimate the spectral index variability intrinsic to the SF galaxies taking into account the measurement errors. Given some intrinsic spectral index distribution, we predict the spectral index distribution that would be observed in the presence of measurement errors. We model the intrinsic spectral index distribution as a Gaussian<sup>b</sup> centred on zero with standard deviation,  $\sigma_{\text{int}}$ . We randomly draw the spectral indices of the 369 SF galaxies included in the spectral index analysis; the spectral index of the  $i^{\text{th}}$  source is given by:

$$\alpha_i = \mathcal{N}(0, \sigma_{\text{int}}) + \mathcal{N}(0, \Delta\alpha_i), \quad (15)$$

<sup>b</sup>It appears reasonable to assume that the intrinsic spectral index distribution has a Gaussian form: the measured distributions of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  for the SF galaxies can be well approximated by Gaussians, and the widths of these distributions are substantially broader than what can be explained by measurement errors alone.



**Figure 18.** Results of Monte Carlo simulations to estimate the variability in  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  intrinsic to the SF galaxies in the GLEAM-6dFGS sample. The dispersion in the spectral index that would be observed given the measurement errors is plotted as a function of the intrinsic spectral index dispersion; the shaded areas show the  $1\sigma$  errors. The dashed line indicates equal values of  $\sigma_{\text{obs}}$  and  $\sigma_{\text{int}}$ . The dotted horizontal lines show the standard deviations of  $\alpha_{\text{low}}$  (0.233) and  $\alpha_{\text{high}}$  (0.148) for the SF galaxies in the real data.

where  $\mathcal{N}(0, \sigma_{\text{int}})$  denotes a random number drawn from a Gaussian distribution of mean 0 and standard deviation  $\sigma_{\text{int}}$ , and  $\Delta\alpha_i$  is the error on the measured spectral index ( $\alpha_{\text{low}}$  or  $\alpha_{\text{high}}$ ) of the  $i^{\text{th}}$  source taken from the GLEAM-6dFGS catalogue.

We then measure the standard deviation of the observed spectral index distribution,  $\sigma_{\text{obs}}$ , as a function of  $\sigma_{\text{int}}$ . We repeat the simulations 1 000 times to improve statistics. Figure 18 shows the mean value of  $\sigma_{\text{obs}}$  as a function of  $\sigma_{\text{int}}$  for  $\alpha_{\text{low}}$  (red curve) and  $\alpha_{\text{high}}$  (blue curve). The observed dispersion in  $\alpha_{\text{low}}$  is higher than that in  $\alpha_{\text{high}}$  due to the larger measurement errors on  $\alpha_{\text{low}}$ . The two curves converge at high  $\sigma_{\text{int}}$  as the measurement errors on  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  become negligible with respect to  $\sigma_{\text{int}}$ .

In the real data, the standard deviations of  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  are 0.233 and 0.148, respectively. From Figure 18, we estimate the intrinsic dispersions in  $\alpha_{\text{low}}$  and  $\alpha_{\text{high}}$  to be  $0.186 \pm 0.012$  and  $0.127 \pm 0.008$ , respectively.